

Surface structure of 45 Hercules: an otherwise unremarkable Ap star with a surprisingly weak magnetic field

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ABSTRACT

The origin of magnetic fields and their role in chemical spot formation on magnetic Ap stars is currently not understood. Here, we contribute to solving this problem with a detailed observational characterization of the surface structure of 45 Her, a weak-field Ap star. We find this object to be a long-period, single-lined spectroscopic binary and determine the binary orbit as well as fundamental and atmospheric parameters of the primary. We study magnetic field topology and chemical spot distribution of 45 Her with the help of the Zeeman Doppler imaging technique. Magnetic mapping reveals the stellar surface field to have a distorted dipolar topology with a surface-averaged field strength of 77 G and a dipolar component strength of 119 G – confirming it as one of the weakest well-characterized Ap-star fields known. Despite its feeble magnetic field, 45 Her shows surface chemical inhomogeneities with abundance contrasts of up to 6 dex. Of the four chemical elements studied, O concentrates at the magnetic equator, whereas Ti, Cr, and Fe avoid this region. Apart from this trend, the positions of Fe-peak element spots show no apparent correlation with the magnetic field geometry. No signs of surface differential rotation or temporal evolution of chemical spots on the time-scale of several years were detected. Our findings demonstrate that chemical spot formation does not require strong magnetic fields to proceed and that both the stellar structure and the global field itself remain stable for sub-100 G field strengths contrary to theoretical predictions.

Key words: stars: atmospheres – stars: chemically peculiar – stars: individual: 45 Her – star: magnetic fields – starspots.

1 INTRODUCTION

About 10 per cent of main sequence A and B-type stars show surface magnetic fields with strengths ranging from a few hundred G up to ~ 30 kG (e.g. Donati & Landstreet 2009; Sikora et al. 2019a). These magnetic fields exhibit a relatively simple surface structure, approximately dipolar in the majority of stars, that remains static on the observable time-scales. Neither the origin nor the incidence of these so-called ‘fossil’ magnetic fields is currently understood (e.g. Neiner et al. 2015). Generally, these fields are believed to be hydrodynamically stable (Braithwaite & Spruit 2004) remnants of one or several magnetic field generation processes operating at an earlier stellar evolutionary stage (Mestel 2003; Moss 2004; Schneider et al. 2019).

The presence of kG-strength magnetic fields in a tenuous atmosphere of an upper main sequence star has a drastic impact on the surface structure and atmospheric physics. Most importantly, various chemical elements separate under the influence of competing processes of radiative levitation and gravitational settling, forming non-uniform horizontal and vertical distributions. This, in turn, gives rise to the observed extreme spectral peculiarities of Ap/Bp (also

known as the upper main sequence chemically peculiar or CP, see Preston 1974) stars and their spot-induced rotational variability. This combination of unique characteristics makes magnetic Ap stars prime targets for investigations of non-trivial interactions between magnetic field and stellar plasma with both observational (e.g. Kochukhov 2017; Kochukhov, Shultz & Neiner 2019 and references therein) and theoretical (e.g. Michaud, Charland & Megessier 1981; Babel & Michaud 1991; Alecian 2015) approaches. These efforts notwithstanding, no successful theoretical model explaining the observed diversity of chemical spot topologies has emerged and the very foundation of the magnetically driven element segregation hypothesis has been challenged by the discovery of time-dependent chemical spot formation in HgMn stars (e.g. Kochukhov et al. 2007, 2011, 2021), which lack organized magnetic fields typical of Ap stars. In this context, it is of particular interest to assess the relationship between magnetic field and chemical spots for well-established magnetic Ap stars at the lower boundary of the range of surface field strength probed by these objects. Are these weak-field Ap stars similar or different compared to stars with stronger fields? Is there any evidence of a systematic reduction of spot contrasts or emergence of temporal variation of spot topologies with the decrease of magnetic field strength? Are the observed field strengths compatible with the theoretically predicted thresholds of fossil field stability?

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The target of the present study – the magnetic Ap star 45 Her (also known as I Her, V776 Her, HR 6234, and HD 151525) – is uniquely suited to address some of these questions. This bright ($V = 5.24$ mag) star is given a B9p CrEu spectral classification in the literature (Renson & Manfroid 2009) and is known to be spectral (Deutsch 1947) and photometric (Burke & Barr 1981) variable with a period of ≈ 4.12 d (e.g. Paurzen et al. 2021). Quantitative modelling of this variability was attempted only once by Hatzes (1991), who derived a surface equivalent width map from a single Cr II spectral line. What sets 45 Her apart from more typical magnetic Ap stars is its illusive magnetic field. Although this star was suspected to be magnetic already by Babcock (1958), the first definitive magnetic field measurement was reported by Aurière et al. (2007). Using relatively low-quality spectropolarimetric data, these authors found a dipolar field strength of $0.54^{+1.38}_{-0.34}$ kG, suggesting that 45 Her might host one of the weakest surface magnetic fields among Ap stars. No information about higher-order magnetic field components or the spatial relationship between the stellar magnetic field topology and chemical spot distributions is available in the literature.

In this study, we present the first detailed magnetic and chemical abundance study of the surface structure of 45 Her based on a large, high-quality spectropolarimetric data set as well as two independent spectroscopic time series collected at different observatories over the time span of 4 yr. This analysis allowed us to gain significant new insights into the surface physics of weak-field Ap stars. The rest of this paper is organized as follows. Section 2 describes observational data employed in our study. Section 3 presents investigation of the spectroscopic binary orbit (Section 3.1), determination of stellar parameters and mean abundances (Section 3.2), derivation of the mean intensity and polarization spectral line profiles (Section 3.3), and their application for surface structure mapping (Section 3.4). The paper is concluded with a summary and discussion of our main findings in Section 4.

2 OBSERVATIONAL DATA

2.1 Narval spectropolarimetry

High-resolution spectropolarimetric observations of 45 Her analysed in this study were obtained with the Narval instrument installed at the 2-m Télescope Bernard Lyot of the Pic du Midi Observatory. Narval is a cross-dispersed échelle, dual-beam spectropolarimeter covering the wavelength range 360–1000 nm with a resolution of $R \approx 65\,000$. Our target was observed 46 times, with two observations acquired in March 2007 followed by an intense monitoring campaign between April and October 2018, which yielded the remaining 44 observations. Circular polarization (Stokes V) spectra were obtained from one to four times a night, using a total exposure time of 1200 s. Each such observation was split into four 300 s subexposures obtained with two different configurations of the polarimetric unit. As described by Donati et al. (1997) and Bagnulo et al. (2009), this dual-beam polarimetric modulation technique allows one to remove spurious instrumental artefacts and derive diagnostic null spectra alongside the intensity and polarization spectra.

The main data reduction and calibration steps were accomplished with an updated version of the ESPRIT code (Donati et al. 1997) executed automatically at the telescope. The reduction products produced by this pipeline are available in the PolarBase¹ archive (Petit et al. 2014). We considered the non-normalized version of

the spectra and performed continuum rectification with the methods described by Rosén et al. (2018). The journal of Narval observations of 45 Her is provided in Table A1, which gives the UT date of each observation, the heliocentric Julian date of mid-exposure, and the signal-to-noise ratio (S/N) in the first three columns. This S/N refers to one pixel, approximately spanning a 1.8 km s^{-1} velocity bin, of the extracted Stokes V spectrum at $\lambda = 550 \text{ nm}$. It was calculated from the standard deviation of the null spectrum and ranges from around 100 to nearly 1000 with a median value of 529.

2.2 CES spectroscopy

Spectroscopic observations of 45 Her were obtained at the TÜBITAK National Observatory (TUG), Antalya, with RTT150 (1.5-metre Russian-Turkish) telescope (Aslan et al. 2001) and Coude Echelle Spectrometer (CES). Observations were carried out on 13 different nights in 2014, 2015, and 2018, with exposure times ranging from 20 to 70 min. In total, 48 spectra of 45 Her with S/N in the 100–200 range were obtained. CES spectra have a resolution of 40 000 and cover a wavelength interval from 369 to 1028 nm. Because of telluric line contamination, low S/N , and continuum normalization problems, only the 400–700 nm region was used in this work.

In each observing night, we obtained 10 biases, five dome flats, and three ThAr (thorium-argon) arc lamp spectra for calibrating the raw spectra of 45 Her. All spectral reduction, wavelength calibration, and spectral normalization step were performed with Image Reduction and Analysis Facility (IRAF; Tody 1986) procedures. The S/N was determined using IRAF by considering the scatter of spectral points in the line-free regions around $\lambda = 550 \text{ nm}$. The average size of CES spectral pixel is 2.4 km s^{-1} .

The journal of CES observations of 45 Her is given in Table A2, which reports the UT and heliocentric Julian dates for each observation, exposure times, achieved S/N , the rotational, and orbital phases.

2.3 Spectrophotometry

Different spectrophotometric observations of 45 Her were collected from the literature for the purpose of modelling the stellar spectral energy distribution (SED) from ultraviolet (UV) to near-infrared. The UV part of the stellar SED is covered by the TD-1 satellite measurements (Thompson et al. 1978). For the 336–1020 nm wavelength region, we used the newly released Gaia DR3 externally calibrated spectrophotometry (Montegriffo et al. 2022). We also considered Geneva photometry (Rufener 1988) converted to absolute fluxes at 7 wavelength bands according to the prescription by Rufener & Nicolet (1988). In the near-infrared, 2MASS photometry (Cutri et al. 2003), as well as JHKLM observations by Groote & Kaufmann (1983), are available. The magnitudes provided by these two catalogues were converted to absolute flux units using the calibrations by Cohen, Wheaton & Megeath (2003) and van der Blik, Manfroid & Bouchet (1996), respectively. The resulting composite observed SED of 45 Her will be discussed and modelled in Section 3.2.

3 ANALYSIS AND RESULTS

3.1 Spectroscopic binary orbit

An initial look at the Narval spectra of 45 Her revealed the presence of $\approx 10 \text{ km s}^{-1}$ peak-to-peak radial velocity changes on a time-scale significantly longer than the stellar rotational period. Although some mentions of possible spectroscopic binarity of 45 Her can be found in the literature (Mason et al. 2001), studies focused on multiplicity

¹ <http://polarbase.irap.omp.eu>

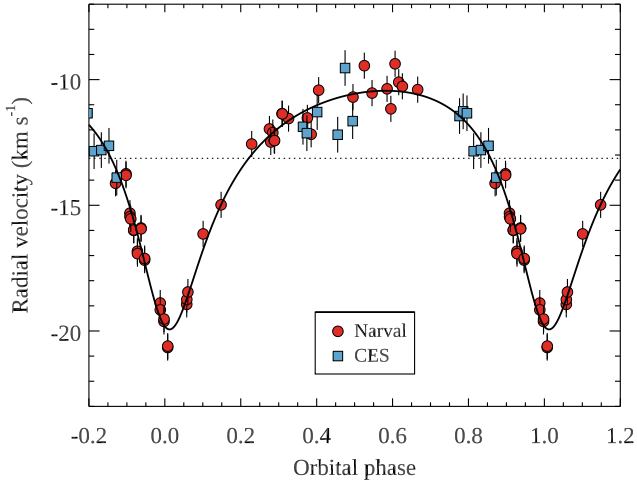


Figure 1. Orbital radial velocity variation of 45 Her. Symbols show the radial velocity measurements derived from the Narval (circles) and CES (squares) observations. The solid line corresponds to the best-fitting orbital solution. The dotted line shows systemic velocity.

Table 1. Spectroscopic orbital parameters and mass function for 45 Her.

Parameter	Value
P_{orb} (d)	99.51 ± 0.10
HJD ₀ (d)	2456702.6 ± 2.1
K_1 (km s ⁻¹)	4.75 ± 0.12
γ (km s ⁻¹)	-13.13 ± 0.09
e	0.445 ± 0.020
ω (°)	167.1 ± 4.2
f (M_{\odot})	$(7.92 \pm 0.67) \times 10^{-4}$

analysis were unable to confirm radial velocity variations and derive orbital parameters (Carrier et al. 2002; Chini et al. 2012). Similar to many other Ap stars, the low-amplitude radial velocity variation of metal lines in 45 Her spectrum is dominated by the rotational variability produced by chemical abundance inhomogeneities. In this case, chemical spots are responsible for the radial velocity modulation between ± 3 and ± 6 km s⁻¹, depending on chemical element. This prevents a straightforward measurement of the orbital radial velocity variation. However, as demonstrated by previous studies (e.g. Bischoff et al. 2017; Kochukhov, Papakonstantinou & Neiner 2022), radial velocity measurements using hydrogen lines are nearly insensitive to heavy element spots and can be successfully employed to investigate the orbital motion of Ap stars.

In this study, we obtained radial velocity measurements from the centre-of-gravity positions of the cores of hydrogen H α to H δ lines. The resulting average radial velocity is reported for Narval spectra in the last column of Table A1. For the lower quality CES spectra, we found that better results are obtained using the core of H α and averaging measurements for each night when multiple spectra were taken. The resulting 13 average radial velocities are given in Table A3.

We fitted a set of SB1 orbital elements (P_{orb} , HJD₀, K_1 , γ , e , ω) to the combined Narval and CES data set using a non-linear least-squares optimization algorithm (Markwardt 2009). Fig. 1 compares the observed radial velocities with the best-fitting orbital solution. The corresponding orbital parameters are given in Table 1 and the orbital phases of Narval observations are reported in the fifth column of Table A1. The scatter of the average hydrogen-line radial velocity

measurements around the orbital solution amounts to 0.52 km s⁻¹ for individual Narval observations and 0.71 km s⁻¹ for nightly averaged CES spectra, which we take as an estimate of the measurement uncertainties for the two data sets.

This analysis establishes 45 Her as the primary in a spectroscopic binary system with P_{orb} of about 100 d, significant eccentricity and a semi-amplitude of only 4.8 km s⁻¹. The radial velocity shifts due to this orbital motion were removed before further spectroscopic and magnetic analysis. The mass function f corresponding to the derived orbital solution implies a minimum mass of the companion $M_2 = 0.193 \pm 0.013 M_{\odot}$ if $M_1 = 2.84 \pm 0.13 M_{\odot}$ (see Section 3.2) is adopted for the primary. If one further assumes the orbital and rotational axes to be aligned ($i_{\text{orb}} = i_{\text{rot}} = 42^\circ$, see Section 3.2), the mass of the secondary is $0.295 \pm 0.018 M_{\odot}$. Thus, the secondary in the 45 Her system is most likely a late-type low-mass star that makes a negligible contribution to the composite high-resolution optical spectrum or a wide-band SED.

3.2 Stellar parameters and mean abundances

Multiple photometric determinations of the atmospheric parameters of 45 Her can be found in the literature (e.g. Kochukhov & Bagnulo 2006; Aurière et al. 2007; Netopil et al. 2008, 2017; Glagolevskij 2019). These analyses reported $T_{\text{eff}} = 9200\text{--}9400$ K and $\log g = 3.5\text{--}3.8$. Based on these results, we carried out an analysis of the mean composition of 45 Her using phase-averaged spectra. We used LLMODELS (Shulyak et al. 2004) model atmosphere with $T_{\text{eff}} = 9250$ K, $\log g = 3.5$, $[M/H] = +0.5$. Abundances were determined using VALD (Ryabchikova et al. 2015) line lists and SYNTH3 (Kochukhov 2007) spectrum synthesis code run from within BINMAG (Kochukhov 2018) fitting and visualization interface. Typically, one to three element abundances together with the projected rotational velocity were determined by fitting theoretical calculations to small segments in the average Narval spectrum of 45 Her. In this analysis, the synthetic spectra were computed under the local thermodynamic equilibrium (LTE) assumption, and therefore lines known to show strong departures from LTE, e.g. the O I near-infrared triplet, were avoided. The resulting abundances are summarized in Table 2. The average projected rotational velocity emerging from this analysis is $v_e \sin i = 36.8 \pm 1.4$ km s⁻¹.

Comparing 45 Her abundances to solar values, we find that this object shows a typical Ap-star abundance pattern (e.g. Ryabchikova 2005) with a moderate underabundance of light elements and significant overabundance of iron-peak elements, particularly Cr. A somewhat less common feature relative to Ap stars with similar T_{eff} is the lack of any Si overabundance and a noticeable depletion of Ti. Our analysis reveals no conclusive evidence of large overabundances of rare-earth or other heavy elements – a feature previously seen in other old Ap stars (Ryabchikova, Leone & Kochukhov 2005; Kochukhov et al. 2006).

Based on this abundance analysis, we computed a custom grid of LLMODELS atmospheres in order to refine T_{eff} and $\log g$. The effective temperature was determined with a global fit to the observed SED (Fig. 2). The stellar angular diameter would normally be another parameter derived in this fit. Considering availability of a relatively accurate Gaia DR3 parallax, 7.72 ± 0.11 mas (Gaia Collaboration 2021), one can replace the angular diameter with the stellar radius in SED modelling. At the same time, results of SED fitting are also influenced by interstellar extinction and reddening. Three-dimensional interstellar medium maps compiled by Lallement et al. (2019) predict $E(B - V) = 0.022 \pm 0.019$, suggesting a

Table 2. Mean abundances of 45 Her compared to the solar composition from Asplund, Amarsi & Grevesse (2021). Uncertain values derived from single spectral features are indicated with “:”.

Element	$\log(N_{\text{el}}/N_{\text{tot}})$	$\log(N_{\text{el}}/N_{\text{tot}})_{\odot}$
He	$<-1.30:$	-1.12
C	-4.24 ± 0.37	-3.57
N	$-5.13:$	-4.20
O	$-4.12:$	-3.34
Mg	-4.70 ± 0.21	-4.48
Al	$-5.48:$	-5.60
Si	-4.61 ± 0.42	-4.52
Ca	-6.62 ± 0.46	-5.73
Ti	-7.69 ± 0.31	-7.06
Cr	-4.71 ± 0.20	-6.41
Mn	-5.63 ± 0.16	-6.61
Fe	-4.16 ± 0.15	-4.57
Ba	-9.40 ± 0.11	-9.76
Y	$-9.67:$	-9.82
Sr	$-9.93:$	-9.20

small and rather uncertain reddening. Owing to availability of the UV spectrophotometry, the colour excess can be independently determined from the SED fit simultaneously with T_{eff} and R . We found that the set of parameters yielding the best description of the observed SED is $T_{\text{eff}} = 9400 \pm 200$ K, $R = 4.3 \pm 0.1 R_{\odot}$, and $E(B - V) = 0.02 \pm 0.01$. The corresponding model SED is compared with the observed fluxes from UV to near-infrared in Fig. 2.

The surface gravity was refined by fitting the $\text{H}\gamma$ line profile in the average Narval spectrum of 45 Her. Theoretical Balmer-line profiles were calculated with the SYNTH3 code for the custom model atmosphere grid described above. As demonstrated by Fig. 3, observations are best reproduced with $\log g = 3.55 \pm 0.05$.

The temperature and radius determined for 45 Her correspond to luminosity $\log L/L_{\odot} = 2.113 \pm 0.042$. Using this information, we examined evolutionary state and estimated the mass of 45 Her by comparing its position in the Hertzsprung–Russell diagram (HRD) with the predictions of MIST² (Choi et al. 2016; Dotter 2016) solar-metallicity theoretical stellar evolutionary tracks. As evident from Fig. 4, 45 Her is an intermediate-mass star approaching the end of its main sequence life. Since evolutionary tracks overlap in this HRD region, one cannot associate a unique mass and age to a given T_{eff} and L pair. Instead, we considered a fine grid of model tracks, calculated with a step of $0.01 M_{\odot}$, and noted all tracks passing within the 1σ error ellipse around 45 Her’s parameters (blue dotted lines in Fig. 4). This procedure yielded masses in the $2.71\text{--}2.97 M_{\odot}$ range and ages between 351 and 469 Myr. The corresponding model tracks have zero-age main sequence radii of $1.77\text{--}1.87 R_{\odot}$.

The summary of atmospheric and physical parameters of 45 Her discussed in this section is given in Table 3.

3.3 Multiline profile analysis

3.3.1 Least-squares deconvolved profiles

Although suspected to be magnetic already by Babcock (1958), 45 Her proved to be one of the most challenging magnetic Ap targets for direct field detection and characterization. Low-resolution investigations (Hubrig et al. 2006; Kochukhov & Bagnulo 2006)

did not detect the field, whereas the high-resolution spectropolarimetric study by Aurière et al. (2007) reported only two definite Zeeman detections out of 14 observations. Polarization signatures are evidently very weak for 45 Her, requiring one to use a multiline technique to enhance the Zeeman signal. Here, we rely on the least-squares deconvolution (LSD; Donati et al. 1997) method as implemented in the iLSD code by Kochukhov, Makaganiuk & Piskunov (2010).

To derive the Stokes I and V LSD profiles, we extracted a custom line list from VALD with the abundances and atmospheric parameters derived in Section 3.2. Spectral regions contaminated by telluric absorption or affected by broad wings of the hydrogen Balmer lines were excluded. Metal lines deeper than 5 per cent of the continuum prior to any macroscopic broadening were selected for the LSD line mask, yielding 2720 line in total. LSD profiles were calculated on the velocity grid spanning a range of $\pm 300 \text{ km s}^{-1}$ using a step of 2 km s^{-1} . The LSD line weights were normalized using the mean wavelength $\lambda_0 = 500 \text{ nm}$ and the effective Landé factor $z_0 = 1.2$ (see Kochukhov et al. 2010 for details).

This application of the LSD procedure to Stokes V observations of 45 Her provided profiles with a median S/N of 15 500, corresponding to a factor 30 gain relative to the original spectra. This allowed us to detect Zeeman signatures at many rotational phases. We used the false alarm probability (FAP) detection diagnostic (Donati, Semel & Rees 1992, Donati et al. 1997) to quantify these detections. Following well-established criteria, we consider a detection to be definite if $\text{FAP} < 10^{-5}$, marginal detection if $10^{-5} < \text{FAP} < 10^{-3}$, and non-detection for $\text{FAP} > 10^{-3}$. This detection diagnostic is reported in the sixth column of Table A1 for all 46 Narval spectra. Considering these original observations without averaging close rotational phases, we get 27 definite detections, 4 marginal detections, and 15 non-detections. Applying the same criteria to the LSD profiles calculated for the diagnostic null spectra yields only non-detections. If 11 pairs of consecutive observations obtained on the same observing nights are averaged, the detection statistics changes to 22 definite detections, one marginal detection, and 12 non-detections out of 35 observations. There are still no significant signals in the null profiles.

In addition to calculating LSD profiles using all available metal lines, we derived several sets of mean profiles for individual chemical elements. Although none of the resulting LSD Stokes V profiles yield usable magnetic signatures, the corresponding LSD Stokes I spectra provide a distilled high S/N observable useful for mapping chemical spot distributions over the stellar surface. Three chemical elements – Fe (1232 lines), Cr (979 lines), and Ti (122 lines) – have sufficiently large number of absorption features in the spectrum of 45 Her for a meaningful application of the LSD method. Average profiles of these chemical elements were derived using the multiprofile mode of the iLSD code, which allowed us to account the background blending by other elements.

3.3.2 Longitudinal magnetic field

The mean longitudinal magnetic field, $\langle B_z \rangle$, is a widely used measure of the disc-averaged line-of-sight component of stellar magnetic geometries (e.g. Mathys 1991). This magnetic characteristic can be determined from the Stokes I and V LSD profiles by calculating the first moment of the V profile about the line’s centre of gravity and normalizing it by the equivalent width of the I profile (Wade et al. 2000; Kochukhov et al. 2010). The results of these calculations for 45 Her using a $\pm 45 \text{ km s}^{-1}$ integration window are presented in the seventh column of Table A1.

²MESA Isochrones & Stellar Tracks, where MESA stands for Modules for Experiments in Stellar Astrophysics (Paxton et al. 2011).

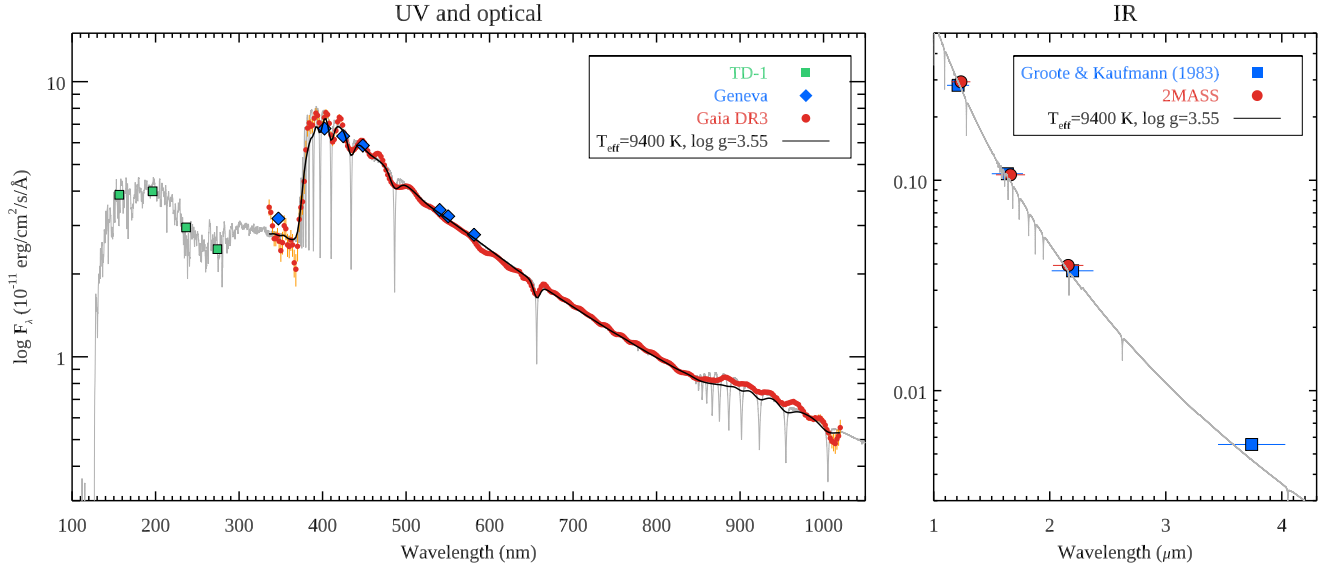


Figure 2. Comparison of the observed and computed spectrophotometry of 45 Her in the UV and optical (left-hand panel) and near-infrared (right-hand panel). Different sources of the observed spectrophotometric data are indicated in the legends. The grey solid line shows the best-fitting theoretical SED for $T_{\text{eff}} = 9400$ K, $\log g = 3.55$, $R = 4.3R_{\odot}$, and $E(B - V) = 0.02$ convolved with a constant Gaussian profile. The black solid line shows the same theoretical model spectrum convolved with a wavelength-dependent broadening function (see Montegriffo et al. 2022) appropriate for the Gaia spectrophotometry.

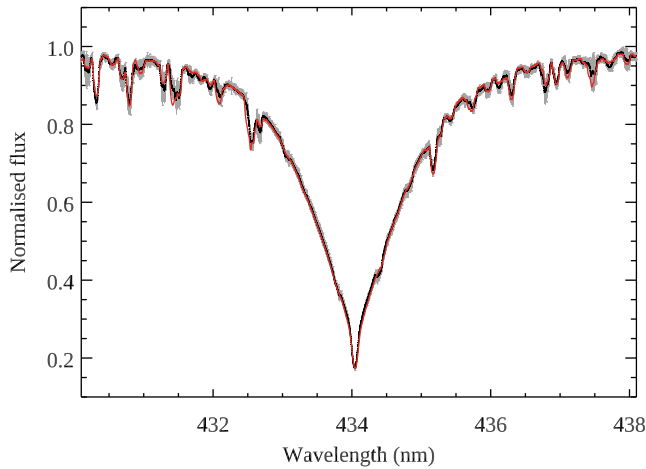


Figure 3. Comparison of the average observed (black symbols) and computed (red solid line) hydrogen $H\gamma$ profile of 45 Her. The grey background curve corresponds to the average spectrum \pm one standard deviation.

This mean longitudinal magnetic field analysis shows that $\langle B_z \rangle$ of 45 Her varies approximately sinusoidally between -100 and $+100$ G. Fig. 5 compares our measurements with the outcome of previous magnetic analysis of this target by Aurière et al. (2007) using the ephemeris derived in Section 3.4.2. It is evident that Narval spectra enable a major improvement in the $\langle B_z \rangle$ measurement precision. Specifically, our $\langle B_z \rangle$ data set has a median error bar of 15 G and yields $>3\sigma$ longitudinal field detections for 24 out of 46 observations. This can be contrasted with a single $>3\sigma$ measurement in Aurière et al. (2007). At the same time, derivation of the mean longitudinal field from the LSD profiles corresponding to the diagnostic null spectra yields no instances of $\langle B_z \rangle$ measurements exceeding 2.3σ , which confirms that our longitudinal field uncertainties are realistic.

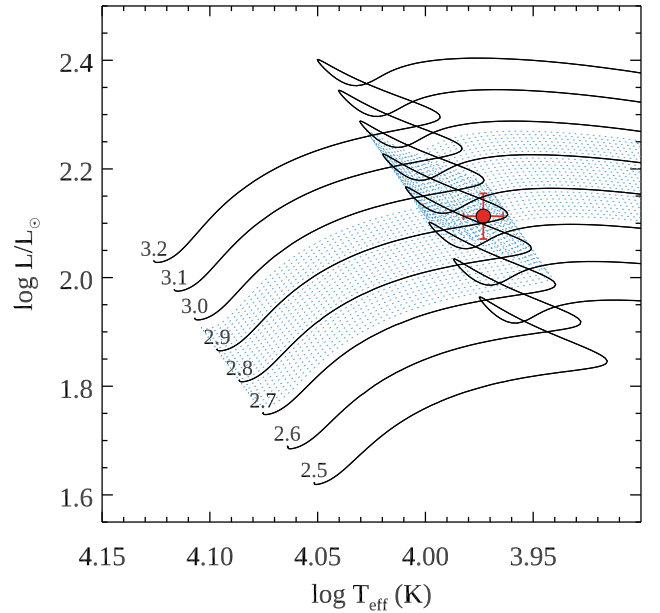


Figure 4. Temperature and luminosity of 45 Her compared to MIST theoretical evolutionary tracks for the initial stellar masses 2.5 – $3.2 M_{\odot}$. Blue dotted lines show the set of evolutionary tracks compatible with T_{eff} and L of 45 Her.

3.4 Zeeman Doppler imaging

3.4.1 Methodology and codes

A variety of inversion codes and tomographic reconstruction methods was employed in our investigation of the surface structure of 45 Her. Most of the inversions were carried out with the INVERSLSD code, introduced by Kochukhov et al. (2014), and subsequently used in several studies of Ap stars (Kochukhov et al. 2017, 2019, 2022). This software enables reconstruction of a vector magnetic

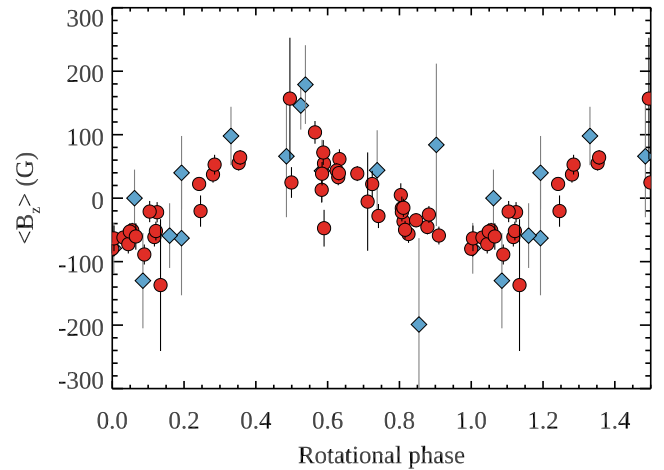
Table 3. Physical parameters of 45 Her.

Parameter	Value
T_{eff} (K)	9400 ± 200
$\log g$ (cm s^{-2})	3.55 ± 0.05
R (R_{\odot})	4.3 ± 0.1
$\log L$ (L_{\odot})	2.113 ± 0.042
M (M_{\odot})	$2.71\text{--}2.97$
Age (Myr)	$351\text{--}469$
P_{rot} (d)	4.116476 ± 0.000022
$v_e \sin i$ (km s^{-1})	35.5 ± 0.5
i ($^{\circ}$)	42.2 ± 1.4

field distribution simultaneously with one scalar parameter map representing local continuum brightness, line strength, temperature, or element abundance. The input observational data for this inversion method comprise time-series of LSD profiles in one or several Stokes parameters. INVERSLSD uses pre-computed tables of local Stokes parameter profiles, which can be calculated under different assumptions and, in particular, treating the polarized radiative transfer problem with different levels of sophistication.

Considering the weakness of the observed circular polarization signatures in the spectrum of 45 Her, we have to rely on modelling the average metal-line LSD Stokes V profiles for the reconstruction of magnetic field topology. The corresponding Stokes I LSD profiles, also obtained by averaging over all metal lines, show strong rotational modulation. This variability corresponds to a weighted average of variabilities of all chemical elements contributing to the LSD line mask (with the largest contributions coming from Fe, Cr, and Ti) and cannot be ascribed to an inhomogeneous surface distribution of one particular element. Nevertheless, it is essential to reproduce this mean metal-line Stokes I LSD profile variability in detail since it affects the Stokes V observable. We model this variability using a non-uniform line strength (or equivalent width) surface distribution implemented in the context of the Unno–Rachkovsky analytical solution of the polarized radiative transfer problem in the Milne–Eddington atmosphere (hereafter ME-UR, Landi Degl’Innocenti & Landolfi 2004). This treatment assumes that the LSD profile behaves as a single spectral line split as a Zeeman triplet with the central wavelength and Landé factor equal to the mean values of these parameters in the LSD mask. This approach to calculating the local Stokes parameter profiles is commonly employed by the ZDI studies of cool active stars (e.g. Morin et al. 2008; Kochukhov & Lavail 2017; Kochukhov & Reiners 2020; Lehtinen et al. 2022), except that here we use the local line strength instead of continuum brightness to model the star-spot signatures in Stokes I spectra.

A different treatment is required to derive physically meaningful surface maps of individual chemical elements from their LSD profiles. As demonstrated by Kochukhov et al. (2010), the response of an LSD profile of particular chemical element to the abundance variation of that element cannot be reproduced by calculations assuming that LSD profile behaves as a single fiducial spectral line. Instead, theoretical LSD profiles correctly reproducing this variation can be computed by applying LSD to grids of detailed radiative transfer calculations for the entire wavelength domain covered by observations. We refer to our previous ZDI studies of Ap stars for an in-depth discussion of the technical implementation and numerical details of this approach (Kochukhov et al. 2014, 2017, 2019, 2022). Here, this methodology is applied to deriving surface distributions of Fe, Cr, and Ti from the Stokes I LSD profiles of these elements

**Figure 5.** Mean longitudinal magnetic field of 45 Her as a function of the rotational phase. Circles correspond to the measurements obtained in this study from Narval spectropolarimetric observations, whereas diamonds show the data from Aurière et al. (2007).

using INVERSLSD; this time with a fixed magnetic field determined from the mean metal-line LSD profiles.

Further chemical elements besides those mentioned above show interesting rotational variability in 45 Her. In particular, modulation of the O I 777.19–777.54 nm triplet is noteworthy. This variability is not amenable to LSD modelling due to the small number of diagnostic lines. In this case, it is more appropriate to apply a conventional DI technique optimized for modelling individual spectral lines. We accomplished this using the INVERS10 ZDI code (Kochukhov & Piskunov 2002; Piskunov & Kochukhov 2002), which has been employed by many previous studies of strong-field Ap stars (e.g. Kochukhov & Wade 2010; Silvester, Kochukhov & Wade 2014a, 2015; Rusomarov et al. 2016; Rusomarov, Kochukhov & Lundin 2018).

Additional complication associated with the spectrum synthesis modelling of the near-infrared O I triplet is a significant departure of the formation of these lines in the atmospheres of A and B-type stars from LTE (e.g. Takeda & Sadakane 1997; Sitnova, Mashonkina & Ryabchikova 2013; Takeda & Honda 2016). Aiming to obtain an accurate O abundance map, we treated departures from LTE with dedicated statistical equilibrium calculations for oxygen atoms in the atmosphere of 45 Her. For this, we used the non-LTE (NLTE) code BALDER (Amarsi et al. 2018a), which is a version of MULTI3D (Leenaarts & Carlsson 2009) with various updates in particular to the equation of state and background opacities. The model atom was the same as that described in Amarsi et al. (2018b) and uses B-spline R-matrix data for electron-impact excitation (Zatsarinny 2006; Tayal & Zatsarinny 2016). This model atom was updated with Stark broadening parameters taken from the Kurucz online data base (Kurucz 1995). The calculations were carried out for a range of O concentrations spanning 6 dex in logarithmic abundance, using a step of 0.2 dex, under the trace-element assumption. INVERS10 was modified to read and interpolate within the resulting tables of NLTE departure coefficients. These data were incorporated in the polarized radiative transfer calculations by modifying the oxygen line opacity and source function according to well-known relations (e.g. Piskunov & Valenti 2017). Fixed magnetic field derived with INVERSLSD from the mean metal-line LSD profiles was included in INVERS10 NLTE O abundance mapping.

In all three tomographic mapping approaches used in this study, the global stellar magnetic field was represented with the help of a spherical harmonic expansion into poloidal and toroidal components (Kochukhov et al. 2014). This expansion included harmonic modes up to $\ell_{\max} = 10$ and was constrained with a harmonic penalty function (Morin et al. 2008; Kochukhov et al. 2014) to minimize contribution of higher-order harmonic terms not justified by the observational data. The line strength and chemical abundance maps were prescribed on a 1876-element surface grid and regularized by minimizing the local surface gradients (Tikhonov regularization, e.g. Piskunov & Kochukhov 2002). The optimal magnetic field and line strength / abundance regularization parameters, corresponding to a compromise between obtaining a satisfactory fit to observations while avoiding spurious small-scale features, were established with the iterative regularization stepping procedure described by Kochukhov (2017).

3.4.2 Rotational period and differential rotation

We started the Doppler imaging analysis of 45 Her by refining the stellar $v_e \sin i$, determining rotational period, and searching for a differential rotation. To this end, we made use of the Stokes *I* LSD profiles derived using all metal lines. Consecutive observations obtained on the same nights were averaged, resulting in a data set comprising 35 spectra. These observations were modelled with a non-uniform surface line strength distribution under the ME-UR local line-profile approach described above. Given the weakness of the stellar magnetic field, it was ignored in this analysis.

Inversions were repeated for different rotational periods around the previously known value $P_{\text{rot}} = 4.1164 \pm 0.0001$ d (Hatzes 1991; Paunzen et al. 2021). Simultaneously, a range of differential rotation behaviours was explored assuming that the stellar surface is sheared according to the simplified solar law commonly adopted by DI and ZDI studies (e.g. Donati et al. 1999; Petit, Donati & Collier Cameron 2002)

$$\Omega(l) = \Omega_{\text{eq}} (1 - \alpha \sin^2 l), \quad (1)$$

where $\Omega(l)$ is the latitudinal dependence of the rotation rate, $\Omega_{\text{eq}} = 1/P_{\text{eq}}$ is the equatorial rotation rate, and $\alpha = \Delta\Omega/\Omega_{\text{eq}}$ is the normalized difference between rotation rates of the equator and pole. Using this approach, the χ^2 of INVERSLSD fit to the LSD profiles was studied as a function of P_{eq} and α with a resolution of 10^{-5} in both of these parameters. This analysis, illustrated in Fig. 6, revealed no evidence of either a solar-like or an antisolar differential rotation with the upper limits 7×10^{-5} (68 per cent confidence) and 11×10^{-5} (90 per cent confidence) derived for the $|\alpha|$ parameter. In the limit of a solid-body rotation, we found $P_{\text{rot}} = P_{\text{eq}} = 4.116476 \pm 0.000022$ d. Adopting the mid-exposure time of the first Narval observation as the reference time yields the ephemeris

$$HJD = 2454173.71327 + 4.116476 \times E, \quad (2)$$

which we use throughout this study.

At this stage of the DI analysis, we also optimized the stellar projected rotational velocity. The best fit to observations was achieved using $v_e \sin i = 35.5 \pm 0.5$ km s $^{-1}$, in good agreement with previous findings (35.0 ± 0.5 km s $^{-1}$ in Hatzes 1991, 35 km s $^{-1}$ in Aurière et al. 2007). Combining P_{rot} and $v_e \sin i$ derived here, with $R = 4.3 \pm 0.1 R_{\odot}$ determined in Section 3.2, one can estimate inclination of the stellar rotational axis, $i = 42.2^\circ \pm 1.4^\circ$, using the relation between the stellar radius, rotational period, and equatorial rotational velocity (Prša et al. 2016). This combination of P_{rot} , $v_e \sin i$, and i

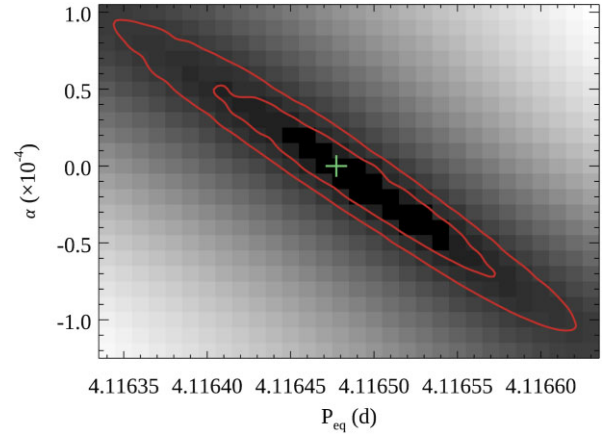


Figure 6. Relative χ^2 of the fit to Stokes *I* LSD profiles as a function of equatorial rotational period P_{eq} and differential rotation parameter α . The solid lines show the 68 per cent and 90 per cent confidence intervals. The cross corresponds to the best-fitting solid-body rotational period.

was employed for all magnetic and abundance inversions reported below. The final comparison of the observed metal-line Stokes *I* LSD profiles with the INVERSLSD fit is shown in the left-hand panel of Fig. 7.

3.4.3 Magnetic field topology

Magnetic field geometry of 45 Her was derived with the ME-UR INVERSLSD modelling of the mean metal-line Stokes *V* profiles after constraining the line strength map as described in the previous section. Three observations with low *S/N* were excluded from this analysis and spectra obtained on the same nights were averaged, resulting in a data set comprising 32 observations. The comparison between the observed and best-fitting model circular polarization spectra is shown in the right-hand panel of Fig. 7. Spherical renderings of the vector magnetic field map corresponding to this fit are presented in Fig. 8.

We find that the global magnetic field of 45 Her is roughly dipolar, featuring a large inclination of the dipolar axis relative to the stellar axis of rotation. At the same time, this dipolar topology is significantly distorted, particularly if one examines the field modulus map (Fig. 8a). There is a noticeable asymmetry between the negative and positive magnetic poles, with the latter one exhibiting stronger field. The local field modulus peaks at 190 G in a spot located within the positive radial field region but does not exceed 166 G in the part of the surface covered by negative field. In both regions, the structures with the strongest field appear to be offset towards the magnetic equator. The surface-averaged field strength is 77 G.

According to the harmonic power spectrum of the derived magnetic field distribution, the field of 45 Her can be characterized (see Table 4) as mostly dipolar (74.6 per cent of the magnetic field energy), poloidal (86.4 per cent), and non-axisymmetric (95.5 per cent). The equivalent dipolar field strength calculated from $\ell = 1$ harmonic coefficients is 119 G. Contribution of $\ell \geq 2$ modes responsible for the apparent distortion of global dipolar field in Fig. 8 is significant. An attempt to fit observations using a purely dipolar, poloidal field increased the standard deviation between the observed and model Stokes *V* spectra by 40 per cent giving an unsatisfactory fit. On the other hand, excluding toroidal components from $\ell_{\max} = 10$ harmonic expansion yields only a marginal decrease of the fit quality while maintaining similar surface magnetic field structure. This suggests

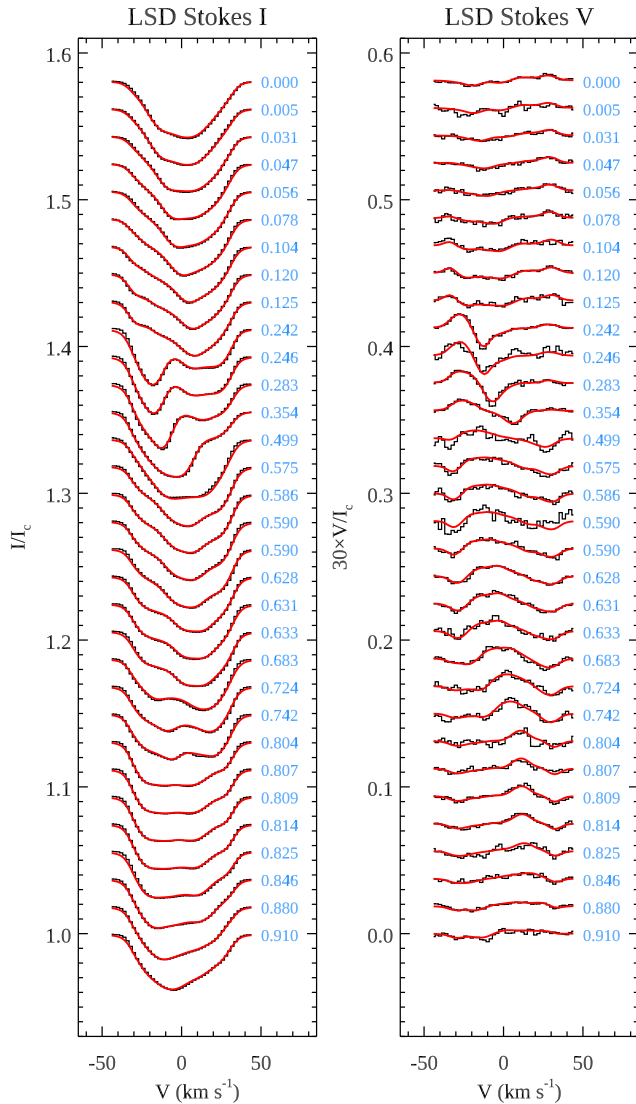


Figure 7. Comparison of the observed (black histograms) and model (red solid lines) LSD Narval Stokes I and V profiles corresponding to the full metal line mask. Profiles are shifted vertically and arranged according to the rotational phase, which is indicated to the right of each profile.

that the toroidal part of 45 Her’s field topology is poorly constrained by the available observations and therefore detection of this field component is rather uncertain.

3.4.4 Chemical spot distributions

Abundance distributions of Ti, Cr, and Fe were derived individually from the Stokes I LSD profiles of these elements using the INVERSLSD code and following the methodology described in Section 3.4.1. The magnetic field was kept fixed in these inversions according to the results of Section 3.4.3. The resulting comparison of the observed and model LSD profiles is illustrated in Figs B1–B5 for Narval data. The same analysis procedure was applied independently to the LSD spectra derived from CES observations. The corresponding model fits are presented in Figs B2–B6. For oxygen, we instead used the INVER10 code modified to account for NLTE effects and directly modelled the near-infrared OI triplet in Narval spectra as illustrated in Fig. B7. This figure also shows the impact of ignoring

departures from LTE, which leads to a major underestimation of the O line absorption for a given oxygen abundance.

The orthographic projection of the O, Ti, Cr, and Fe spot distributions obtained from Narval observations is presented in Fig. 9, where the star is shown at five rotational phases and at the actual inclination angle as seen by the observer. The same plot for Ti, Cr, and Fe spot maps derived from CES data is given in Fig. 10. A comparison between these two sets of chemical abundance maps will be discussed in Section 4.4.

For all four studied chemical elements, we found large-scale, high-contrast abundance inhomogeneities spanning several dex in logarithmic element concentration. The O map exhibits a ring-like structure with a contrast reaching nearly 6 dex. Additionally, there is a spot characterized by an extreme concentration of O at one of the two intersections of this ring with the stellar rotational equator. This asymmetry of the reconstructed O distribution is not an inversion artefact since the O triplet lines clearly appear stronger and wider at phase ≈ 0.35 compared to phase ≈ 0.80 when the two sides of the overabundance ring cross the disc centre. The O map reconstructed in LTE is qualitatively similar but has much higher local abundance values.

The Ti distribution exhibits the second-highest surface contrast of around 5 dex. The behaviour of this element resembles that of O, although in addition to a ring-like structure it shows a number of smaller spots and a more articulated high-contrast feature at phase 0.3. The location of this spot is close to the O overabundance region at the stellar equator.

In contrast to O and Ti, the distribution of Cr is dominated by a ring of relative underabundance spanning the visible hemisphere. For this element, our inversion procedure also recovers major overabundance features below the rotational equator. However, as seen from Fig. 9, these parts of the stellar surface are always visible near the limb and contribute little to the disc-integrated spectra. Reliability of these surface features is therefore low.

Finally, Fe shows the lowest abundance contrast (about 1.5 dex) and the most complex surface distribution among all four studied chemical elements. The iron map exhibits at least five latitudinally extended overabundance spots or arcs. This complexity can be traced back to the rotational modulation of Fe lines (e.g. Fig. B5), which shows many small-amplitude distortions travelling across rotationally broadened profile compared to a couple of dominant spot features visible in the lines of O, Ti, and Cr.

4 SUMMARY AND DISCUSSION

In this study, we presented a detailed investigation of the physical parameters, binarity, magnetic field, and surface chemical spot distributions of the magnetic Ap star 45 Her. This star was found to be a wide, single-lined spectroscopic binary with $P_{\text{orb}} = 99.7$ d and a radial velocity semi-amplitude of $K_1 = 4.7$ km s $^{-1}$. Its companion is a late-type, low-mass star that makes no detectable contribution to the composite spectra. We revised atmospheric parameters of 45 Her based on model atmosphere fitting of the H γ line profile and several spectrophotometric observations, including recently released Gaia DR3 spectrophotometry. These data are best reproduced with $T_{\text{eff}} = 9400$ K, $\log g = 3.55$, and $R = 4.3 R_{\odot}$. Abundance analysis based on the mean spectra revealed a typical chemical pattern of an evolved magnetic Ap star (Kochukhov et al. 2006; Bailey, Landstreet & Bagnulo 2014): an underabundance of light elements, Ca and Ti, approximately normal concentration of Mg and Si, and large overabundance of Fe-peak elements without a detectable enhancement of rare-earths. Zeeman circular polarization signatures were repeatedly detected in high-resolution spectropolarimetric observations

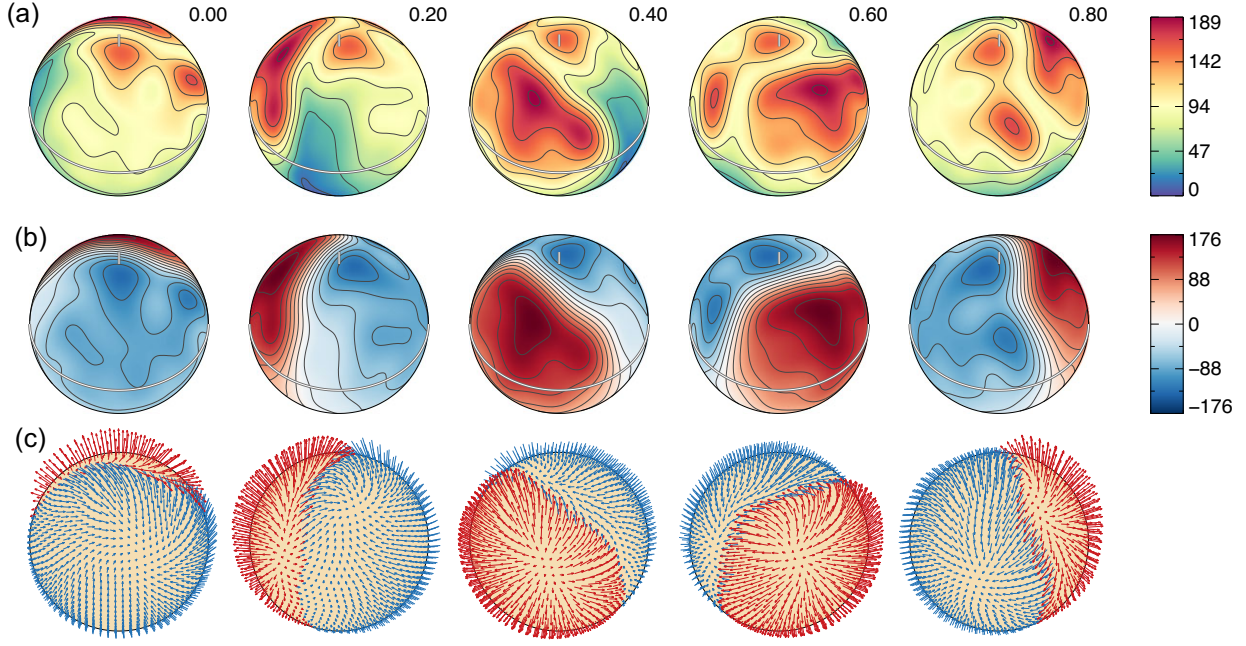


Figure 8. Magnetic field topology of 45 Her derived with ZDI analysis from metal-line Stokes V LSD profiles. The three rows of spherical plots show maps of the magnetic field strength (a), the radial field component (b), and the magnetic field vector (c). The star is shown at the inclination $i = 42^\circ$ for five rotational phases, which are indicated next to each column. The contours over spherical maps in panels (a) and (b) are plotted with a step of 25 G. The stellar equator and visible rotational pole are indicated with double lines. The side colour bars give the field strength in G.

Table 4. Magnetic field characteristics of 45 Her.

Parameter	Value
E_{pol}	86.4 per cent
$E(lm_l > 0)$	95.5 per cent
$E(\ell = 1)$	74.6 per cent
$E(\ell = 2)$	15.3 per cent
$E(\ell = 3)$	4.3 per cent
$E(\ell \geq 4)$	5.8 per cent
$\langle B \rangle$	77.5 G
$\langle B_r \rangle$	66.1 G
B_d	118.8 G

of 45 Her, indicating a reversing longitudinal magnetic field with a strength below 100 G. We used the LSD circular polarization profiles to reconstruct the surface magnetic field topology of 45 Her with the help of the ZDI technique. The spatially resolved surface magnetic field strength of that star peaks at 190 G with the mean field modulus of 77 G. The stellar field geometry is dominated by the poloidal, non-axisymmetric components and is roughly dipolar, although contributions of higher-order components (quadrupole and above) is non-negligible. Applying Doppler imaging analysis to LSD intensity profiles derived from all metal lines, we demonstrated the absence of a latitudinal differential rotation with an upper limit of $|\alpha| = |\Delta\Omega/\Omega| \approx 10^{-4}$. We then used the LSD intensity profiles of Ti, Cr, and Fe to reconstruct spot distributions of these elements. This analysis was repeated for two independent data sets obtained at different observatories. Finally, we used the higher resolution and better S/N Narval data set to recover the oxygen surface abundance map from the spectacular variation of the near-infrared O I triplet based on the spectrum synthesis calculations that incorporated NLTE line formation of that element.

4.1 Magnetic field of 45 Her compared to other Ap stars

The distinguishing feature of 45 Her is its remarkably weak surface field for a magnetic Ap star. These objects are known to have fields with dipolar components ranging from several hundred G up to ~ 30 kG (e.g. Aurière et al. 2007; Sikora et al. 2019b). However, the accuracy and reliability of these field strength estimates, especially for weaker fields, depend on the quality of observations and degree of detail with which these fields are studied. To this end, 45 Her has the lowest magnetic field strength among Ap stars modelled with ZDI (Kochukhov et al. 2019) and the second-lowest field among all Ap stars analysed with modern high-resolution spectropolarimetry. The only star with a weaker field, HD 5550 (Alecian et al. 2016), is an atypical object in several respects. It is a short-period ($P_{\text{orb}} = 6.84$ d), double-lined spectroscopic binary. The field with a dipolar strength of $B_d = 65 \pm 20$ G is nearly aligned with the orbital axis and is found on the primary component, which rotates with the same period as its orbital motion and lacks significant spectral variations indicative of chemical spots. In contrast, 45 Her is effectively a single object that exhibits all staple characteristics of magnetic Ap stars (non-solar chemical abundances, highly inclined distorted dipolar field, prominent rotational modulation due to chemical spots). This proves that the chemical peculiarity and surface structure formation phenomena associated with a fossil magnetic field in late-B and A-type stars extend down to dipolar field strengths of ≈ 120 G (corresponding to a surface-averaged field strength of ≈ 80 G). This is lower than thought previously but still noticeably higher than the ≈ 30 G dipolar field of the Am star γ Gem (Blazère et al. 2020) and orders of magnitude stronger than sub-G magnetic fields found on other Am and normal intermediate-mass stars (Petit et al. 2010, 2011; Blazère et al. 2016).

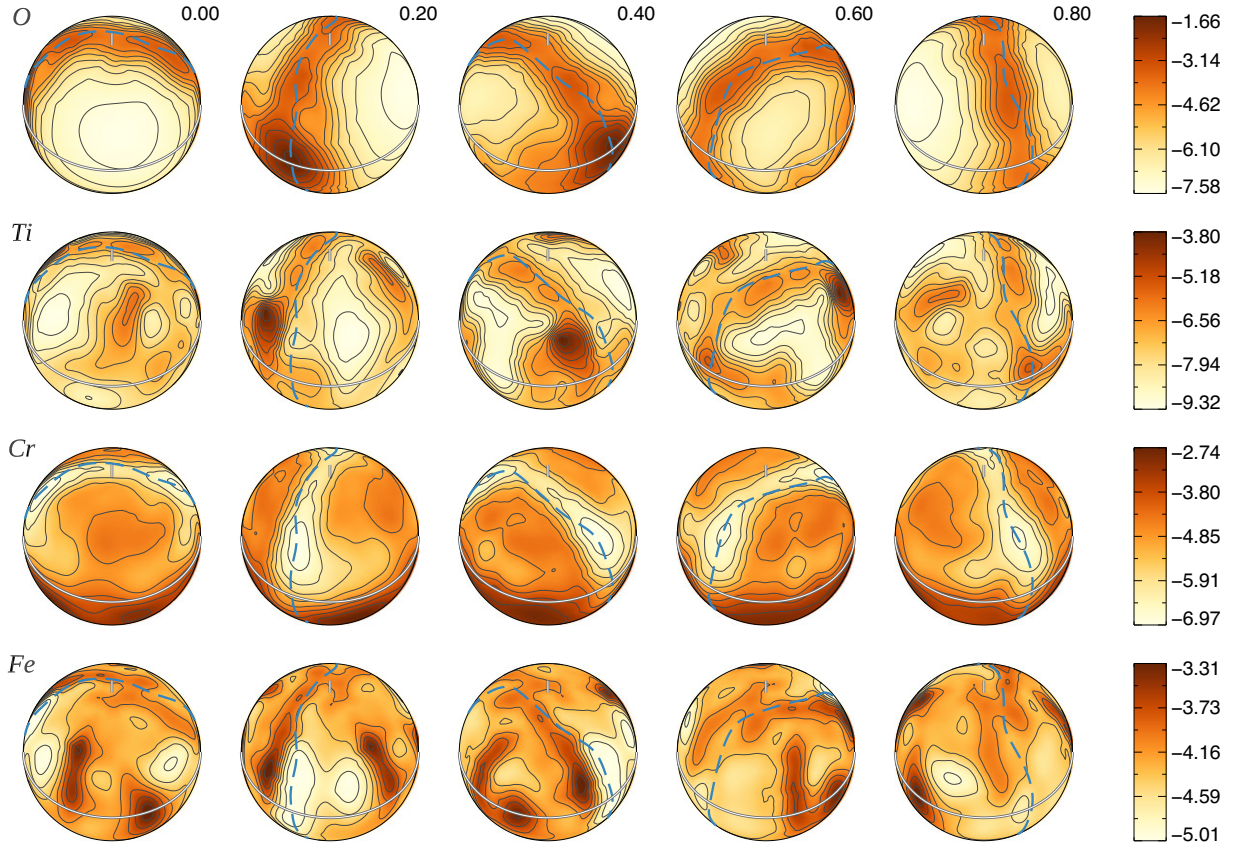


Figure 9. Surface abundance distributions of O, Ti, Cr, and Fe derived with DI modelling of Stokes *I* Narval spectra obtained during 2018. The star is shown at the inclination $i = 42^\circ$ for five rotational phases, which are indicated next to each column. The contours over spherical maps are plotted with a step of 0.25 dex for Fe and 0.50 dex for other elements. The fractional element abundances are given by the side bars in the $\log(N_{e1}/N_{\text{tot}})$ units. The double lines correspond to the stellar equator and rotational pole. The thick dashed line shows where the radial magnetic field changes sign according to the magnetic map in Fig. 8.

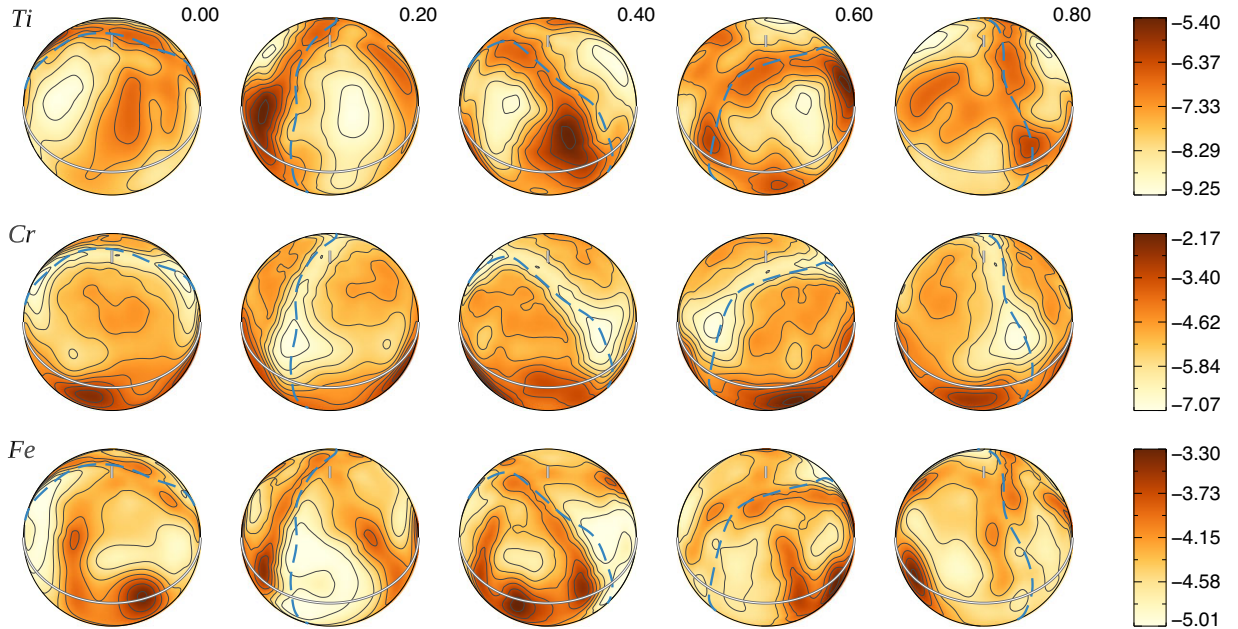


Figure 10. Same as Fig. 9 but for the surface abundance distributions of Ti, Cr, and Fe derived from CES observations obtained between 2014 and 2018.

4.2 Stability of weak fossil magnetic fields

Owing to the weakness of its magnetic field, 45 Her is a key object in the discussion of the hydromagnetic stability of fossil fields in early-type stars and of the ensuing lower field strength threshold. Based on a sample of 24 weak-field magnetic Ap stars, Aurière et al. (2007) has established that these objects appear to possess dipolar fields stronger than 300 G with the median strength estimated to be $B_d \approx 1.4$ kG. The existence of this 300 G threshold field strength was confirmed by Sikora et al. (2019b), who studied a volume-limited sample of 52 magnetic Ap stars with higher quality data. They concluded that the mean dipolar field strength of their sample of Ap stars was 2.6 kG and also could not find a single star with B_d unambiguously below the 300 G threshold.

To explain the apparent lower limit of Ap-star fields, Aurière et al. (2007) put forward a hypothesis that the global fossil field is subject to a Taylor-type instability (e.g. Spruit 1999) and cannot survive below a certain minimum field strength B_c , which depends on the atmospheric equipartition field B_{eq} , the stellar rotational period, radius, and effective temperature,

$$B_c = 2B_{eq} \left(\frac{P_{rot}}{5d} \right)^{-1} \left(\frac{R}{3R_\odot} \right) \left(\frac{T_{eff}}{10^4 K} \right)^{-1/2}. \quad (3)$$

All stars studied by Aurière et al. (2007) and Sikora et al. (2019b) satisfied the condition $B_d \geq B_c$ but this is not the case for 45 Her. Its equipartition field is $B_{eq} = 131$ G at $\tau_{500} = 2/3$ for the $T_{eff} = 9400$ K, $\log g = 3.55$ model atmosphere calculated in our study. Inserting other relevant stellar parameters in equation (3) yields $B_c = 471$ G. This is about a factor of 4 higher than the observed dipolar field strength of 119 G. Thus, the field of 45 Her is clearly incompatible with the theory underlying the commonly discussed ≈ 300 G dipolar strength limit. Similar conclusions can be reached for the primary component of HD 5550 (Alecian et al. 2016) with the already mentioned caveat that the latter is part of a close binary system, which is likely to be relevant for both the origin and long-term survival of the fossil field. In contrast, a wide orbit and low mass of 45 Her's binary companion allows one to treat the primary as effectively a single star in the context of the field stability discussion.

Another type of field stability assessment is commonly carried out in the context of 1D modelling of subsurface convection zones in the intermediate- and high-mass stars (e.g. Cantiello et al. 2009; Cantiello & Braithwaite 2011, 2019). This treatment explores stability of the hydrogen and helium ionization zones in lower mass stars and Fe-peak opacity zone in higher mass stars as a function of vertical magnetic field strength. The underlying idea is that the field strength has to exceed a certain limit, B_c , to provide stability to convection. Fields weaker than B_c will be dragged, tangled, and quickly destroyed by convective motions. Arguing along these lines, Jermyn & Cantiello (2020) and Farrell et al. (2022) applied the convective stability criterium by MacDonald & Petit (2019; originally developed for O-star envelopes) to grids of 1D stellar structure models covering the mass interval relevant for magnetic Ap stars. According to the latest results by Farrell et al. (2022), the fundamental parameters of 45 Her require B_c of at least 1500 G. In this theory, the critical field corresponds to the surface-averaged radial magnetic field component. But, according to our ZDI results, the latter amounts to just 66 G for 45 Her. Thus, there is a major contradiction between these theoretical predictions and observations. This tension points to a problem with the concept of a minimal vertical field required for stability to convection, likely stemming from an approximate 1D treatment of convection in the current stellar structure models. This also means that the far-reaching conclusions, for example on the

separation between fossil and dynamo fields (Jermyn & Cantiello 2020) and on the initial magnetic field distribution of AB stars (Farrell et al. 2022), obtained using this theoretical framework have to be revised to account for the existence of stars like 45 Her and HD 5550.

4.3 Relation between magnetic field and chemical spots

We now take a closer look at the relation between magnetic field geometry and locations of chemical spots using the visualization procedure introduced by Kochukhov et al. (2022). Fig. 11 shows the Hammer–Aitoff maps of the radial magnetic field component, the field modulus, and the local field inclination relative to the surface normal simultaneously with the abundance distributions of O, Ti, Cr, and Fe. The parts of the stellar surface completely or partially invisible (as defined in Kochukhov et al. 2022) to the observer are indicated with shaded regions.

Starting with oxygen, we see a clear correlation of that element with the horizontal field regions. The O overabundance forms nearly a perfect ring at the magnetic equator, similar to the behaviour of this element in a few other magnetic Ap stars (Rice, Wehlau & Holmgren 1997; Kochukhov et al. 2004). In 45 Her, this symmetry is, however, broken by an additional O spot located at one of the intersections of the magnetic and rotational equators. Turning to the field strength map, we can see that at this location, the field intensity is systematically lower compared to the other side of the magnetic equator.

The distribution of Ti features a major spot that is located not far from the O spot, but is offset from the magnetic equator. There are two other, less prominent, Ti spots with the first one located at another intersection of the magnetic and rotational equators and the second one roughly coinciding with the positive magnetic pole. Titanium was rarely targeted by DI analyses of Ap stars, and a few previous maps of this element reconstructed simultaneously with magnetic field did not show the tendency of concentration at the magnetic equator observed in 45 Her (Lüftinger et al. 2010; Silvester, Kochukhov & Wade 2014b; Rusomarov et al. 2016).

The surface distribution of Cr is dominated by a relative underabundance ring at the magnetic equator. This result agrees the earlier finding for 45 Her and two other Ap stars by Hatzes (1991). However, the trend for Cr to exhibit a depletion at the magnetic equator is rarely observed in Ap stars analysed with DI (e.g. Kochukhov et al. 2017; Kochukhov et al. 2019; Kochukhov et al. 2022).

The surface map of Fe reconstructed for 45 Her is the most complex of all studied elements, comprising at least five separate Fe overabundance spots and three regions of prominent Fe depletion. One of the overabundance spots is co-located with the previously discussed major Ti spot in the vicinity of the weaker-field segment of the magnetic equator. At the same time, Fe is underabundant at this part of the magnetic equator itself, similar to Cr. Other features that can be discerned in the Fe surface map are pairs of spots located close to both negative and positive magnetic poles. Overall, the behaviour of Fe in 45 Her shares the trend, seen in the DI results for other Ap stars (e.g. Kochukhov et al. 2004, 2014; Kochukhov et al. 2019), for this element to exhibit a complex multispot distribution lacking a direct relation to the magnetic field geometry.

4.4 Long-term stability of abundance distributions

Independent abundance inversions using Narval and CES observations offer a unique opportunity to probe possible long-term evolution of chemical spots or, alternatively, assess uncertainties of

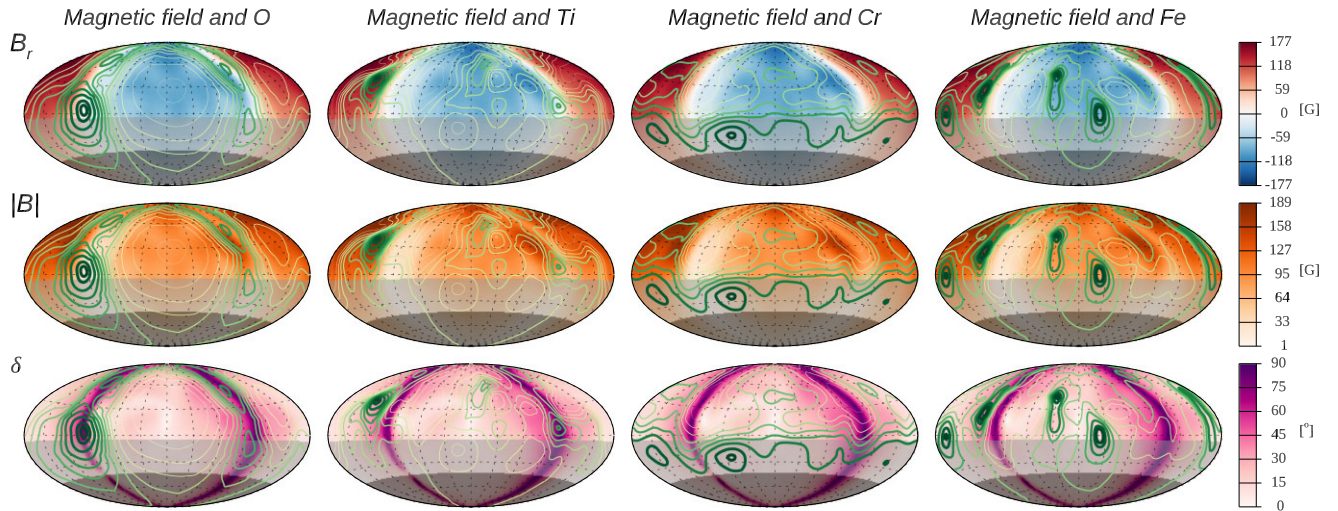


Figure 11. Comparison of the magnetic field topology of 45 Her with abundance distributions for O, Ti, Cr, and Fe based on Narval observations. The three rows of Hammer–Aitoff projection plots display in the background, from top to bottom, the radial magnetic field component, the field strength, and the local field inclination. The dotted lines show the longitude–latitude grid plotted with a 30° step and zero longitude in the centre. The shaded bottom part of the maps indicates latitudes that are completely invisible (dark grey) or have a relative visibility of less than 25 per cent (light grey). Chemical abundance distributions are shown with foreground contours plotted every 0.25 dex for Fe and 0.50 dex for other elements. The darker and thicker contours correspond to regions with a higher element abundance.

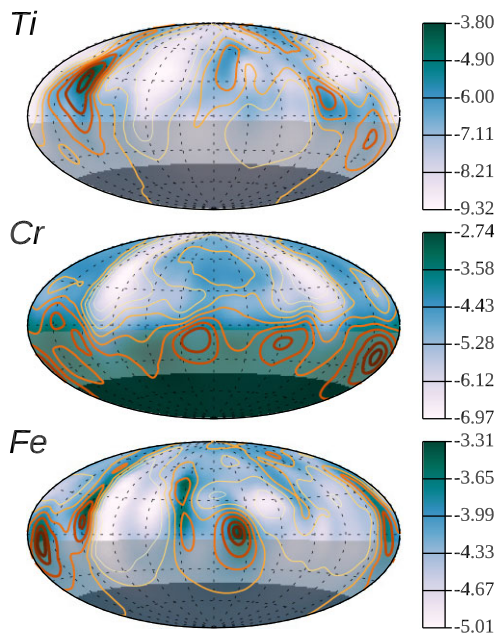


Figure 12. Agreement between abundance distributions of Ti, Cr, and Fe reconstructed from independent spectroscopic data sets. The background Hammer–Aitoff plots show the abundance maps derived from the Narval spectra. The foreground contours correspond to the chemical spot distributions obtained from the CES observations. The contours are plotted every 0.25 dex for Fe and 0.50 dex for other elements, with darker colour and thicker lines indicating a higher element abundance.

DI reconstruction assuming a static surface distribution. It is known that, in the absence of strong global magnetic fields similar to those usually found in Ap stars, heavy-element spots on HgMn stars evolve on the time-scales from years (Kochukhov et al. 2007) to months (Korhonen et al. 2013). Similar behaviour cannot be a priori excluded for 45 Her, considering the unusually weak magnetic field of this star.

All but two Narval observations were obtained in 2018, whereas most of the CES spectra were collected in 2014–2015. There is, therefore, a 3-yr gap between the bulk of the two data sets. The spectral resolution, typical S/N , rotational phase coverage, and reduction procedures applied to the Narval and CES data were all different. Nevertheless, the comparison of Ti, Cr, and Fe maps in Fig. 12 shows an excellent agreement between the outcomes of the DI inversions using the two sets of spectra. In particular, every spot feature discussed above for the surface maps based on the Narval observations is reproduced in the maps inferred from the CES data. This even includes the most complex Fe map, for which the two independent inversions yield the same five overabundance spots. The only systematic difference between the two sets of maps, most clearly seen by comparing Figs 9 and 10, is a smoother surface distribution derived using the CES observations, likely reflecting a lower resolution and S/N of that data set. Thus, there is no evidence of any change in the surface maps of these three chemical elements between 2014–2015 and 2018. In this respect, 45 Her behaves similar to other magnetic Ap stars and differently from non-magnetic HgMn stars.

Assuming static surface abundance distributions, we computed a mean absolute difference (weighted by the surface element visibility) of 0.32, 0.22, and 0.12 dex for Ti, Cr, and Fe, respectively. These numbers can be considered as conservative estimates of the local uncertainty of abundance reconstructions in this study.

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DATA AVAILABILITY

The spectropolarimetric Narval observations underlying this article are available from the PolarBase archive (<http://polarbase.irap.omp.eu>). The CES observations will be shared on reasonable request to the corresponding author.

REFERENCES

- AlecianG., 2015, *MNRAS*, 454, 3143
 AlecianE., TkachenkoA., NeinerC., FolsomC. P., LeroyB., 2016, *A&A*, 589, A47
 AmarsiA. M., NordlanderT., BarklemP. S., AsplundM., ColletR., LindK., 2018a, *A&A*, 615, A139
 AmarsiA. M., BarklemP. S., AsplundM., ColletR., ZatsarinnyO., 2018b, *A&A*, 616, A89
 AslanZ. et al., 2001, *Astron. Lett.*, 27, 398
 AsplundM., AmarsiA. M., GrevesseN., 2021, *A&A*, 653, A141
 AurièreM. et al., 2007, *A&A*, 475, 1053
 BabcockH. W., 1958, *ApJ*, 128, 228
 BabelJ., MichaudG., 1991, *ApJ*, 366, 560
 BagnuloS., LandolfiM., LandstreetJ. D., Landi Degl'InnocentiE., FossatiL., SterzikM., 2009, *PASP*, 121, 993
 BaileyJ. D., LandstreetJ. D., BagnuloS., 2014, *A&A*, 561, A147
 BischoffR. et al., 2017, *Astron. Nachr.*, 338, 671
 BlazèreA. et al., 2016, *A&A*, 586, A97
 BlazèreA., PetitP., NeinerC., FolsomC., KochukhovO., MathisS., DealM., LandstreetJ., 2020, *MNRAS*, 492, 5794
 BraithwaiteJ., SpruitH. C., 2004, *Nature*, 431, 819
 BurkeE. W. J., BarrT. H., 1981, *PASP*, 93, 344
 CantielloM., BraithwaiteJ., 2011, *A&A*, 534, A140
 CantielloM., BraithwaiteJ., 2019, *ApJ*, 883, 106
 CantielloM. et al., 2009, *A&A*, 499, 279
 CarrierF., NorthP., UdryS., BabelJ., 2002, *A&A*, 394, 151
 ChiniR., HoffmeisterV. H., NasseriaA., StahlO., ZinneckerH., 2012, *MNRAS*, 424, 1925
 ChoiJ., DotterA., ConroyC., CantielloM., PaxtonB., JohnsonB. D., 2016, *ApJ*, 823, 102
 CohenM., WheatonW. A., MegeathS. T., 2003, *AJ*, 126, 1090
 CutriR. M. et al., 2003, VizieR Online Data Catalog. p. II/246
 DeutschA. J., 1947, *ApJ*, 105, 283
 DonatiJ.-F., LandstreetJ. D., 2009, *ARA&A*, 47, 333
 DonatiJ.-F., SemelM., ReesD. E., 1992, *A&A*, 265, 669
 DonatiJ. F., SemelM., CarterB. D., ReesD. E., Collier CameronA., 1997, *MNRAS*, 291, 658
 DonatiJ.-F., Collier CameronA., HussainG. A. J., SemelM., 1999, *MNRAS*, 302, 437
 DotterA., 2016, *ApJS*, 222, 8
 FarrelleE., JermynA. S., CantielloM., Foreman-MackeyD., 2022, *ApJ*, 938, 10
 Gaia Collaboration, 2021, *A&A*, 649, A1
 GlagolevskijY. V., 2019, *Astrophys. Bull.*, 74, 66
 GrooteD., KaufmannJ. P., 1983, *A&AS*, 53, 91
 HatzesA. P., 1991, *MNRAS*, 253, 89
 HubrigS., NorthP., SchöllerM., MathysG., 2006, *Astron. Nachr.*, 327, 289
 JermynA. S., CantielloM., 2020, *ApJ*, 900, 113
 KochukhovO., 2007, in RomanyukI. I., KudryavtsevD. O., eds, Physics of Magnetic Stars. SAO RAS, Niznij Arkhyz, p. 109
 KochukhovO., 2017, *A&A*, 597, A58
 KochukhovO., 2018, Astrophysics Source Code Library, record ascl:1805.015
 KochukhovO., BagnuloS., 2006, *A&A*, 450, 763
 KochukhovO., LavailA., 2017, *ApJ*, 835, L4
 KochukhovO., PiskunovN., 2002, *A&A*, 388, 868
 KochukhovO., ReinersA., 2020, *ApJ*, 902, 43
 KochukhovO., WadeG. A., 2010, *A&A*, 513, A13
 KochukhovO., DrakeN. A., PiskunovN., de la RezaR., 2004, *A&A*, 424, 935
 KochukhovO., TsybmalV., RyabchikovaT., MakaganykV., BagnuloS., 2006, *A&A*, 460, 831
 KochukhovO., AdelmanS. J., GulliverA. F., PiskunovN., 2007, *Nature Phys.*, 3, 526
 KochukhovO., MakaganiukV., PiskunovN., 2010, *A&A*, 524, A5
 KochukhovO. et al., 2011, *A&A*, 534, L13
 KochukhovO., LüftingerT., NeinerC., AlecianE., MiMeS Collaboration, 2014, *A&A*, 565, A83
 KochukhovO., SilvesterJ., BaileyJ. D., LandstreetJ. D., WadeG. A., 2017, *A&A*, 605, A13
 KochukhovO., ShultzM., NeinerC., 2019, *A&A*, 621, A47
 KochukhovO., KhalackV., KobzarO., NeinerC., PaunzenE., Labadie-BartzJ., David-UrazA., 2021, *MNRAS*, 506, 5328
 KochukhovO., PapakonstantinouN., NeinerC., 2022, *MNRAS*, 510, 5821
 KorhonenH. et al., 2013, *A&A*, 553, A27
 KuruczR. L., 1995, in AdelmanS. J., WieseW. L., eds, ASP Conf. Ser. Vol. 78, Astrophysical Applications of Powerful New Databases. Astron. Soc. Pac., San Francisco, P. 205
 LallentR., BabusiauxC., VergelyJ. L., KatzD., ArenouF., ValetteB., Hot-tierC., Capitaniol., 2019, *A&A*, 625, A135
 Landi Degl'InnocentiE., LandolfiM., 2004, Polarization in Spectral Lines. Astrophysics and Space Science Library, Vol. 307, Kluwer Academic Publishers, Dordrecht
 LeenaartsJ., CarlssonM., 2009, in LitesB., CheungM., MagaraT., MariskaJ., ReevesK., eds, ASP Conf. Ser. Vol. 415, The Second Hinode Science Meeting: Beyond Discovery-Toward Understanding. Astron. Soc. Pac., San Francisco, p. 87
 LehtinenJ. J. et al., 2022, *A&A*, 660, A141
 LüftingerT. et al., 2010, *A&A*, 509, A43
 MacDonaldJ., PetitV., 2019, *MNRAS*, 487, 3904
 MarkwardtC. B., 2009, in BohlenderD. A., DurandD., DowlerP., eds, ASP Conf. Ser. Vol. 411, Astronomical Data Analysis Software and Systems XVIII. Astron. Soc. Pac., San Francisco, p. 251
 MasonB. D., WycoffG. L., HartkopfW. I., DouglassG. G., WorleyC. E., 2001, *AJ*, 122, 3466
 MathysG., 1991, *A&AS*, 89, 121
 MestelL., 2003, in BalonaL. A., HenrichsH. F., MedupeR., eds, ASP Conf. Ser. Vol. 305, Magnetic Fields in O, B, and A Stars: Origin and Connection to Pulsation, Rotation and Mass Loss. Astron. Soc. Pac., San Francisco, p. 3
 MichaudG., CharlandY., MegessierC., 1981, *A&A*, 103, 244
 MontegriffoP. et al., 2022, preprint([arXiv:2206.06205](https://arxiv.org/abs/2206.06205))
 MorinJ. et al., 2008, *MNRAS*, 390, 567
 MossD., 2004, in ZverkoJ., ZiznovskyJ., AdelmanS. J., WeissW. W., eds, Proc. IAU Symp. Vol. 224, The A-Star Puzzle. Cambridge Univ. Press, Cambridge, p. 245
 NeinerC., MathisS., AlecianE., EmeriauC., GrunhutJ., BinaMiCS, MiMeS Collaborations, 2015, in NagendraK. N., BagnuloS., CentenoR., Jesús Martínez GonzálezM., eds, Proc. IAU Symp. Vol. 305, Polarimetry. Cambridge Univ. Press, Cambridge, p. 61
 NetopilM., PaunzenE., MaitzenH. M., NorthP., HubrigS., 2008, *A&A*, 491, 545
 NetopilM., PaunzenE., HümmerichS., BernhardK., 2017, *MNRAS*, 468, 2745
 PaunzenE., SupíkováJ., BernhardK., HümmerichS., PrišegenM., 2021, *MNRAS*, 504, 3758
 PaxtonB., BildstenL., DotterA., HerwigF., LesaffreP., TimmesF., 2011, *ApJS*, 192, 3
 PetitP., DonatiJ. F., Collier CameronA., 2002, *MNRAS*, 334, 374

- PetitP. et al., 2010, *A&A*, 523, A41
- PetitP. et al., 2011, *A&A*, 532, L13
- PetitP., LougeT., ThéadoS., PaletouF., MansetN., MorinJ., MarsdenS. C., JeffersS. V., 2014, *PASP*, 126, 469
- PiskunovN., KochukhovO., 2002, *A&A*, 381, 736
- PiskunovN., ValentiJ. A., 2017, *A&A*, 597, A16
- PrestonG. W., 1974, *ARA&A*, 12, 257
- PršaA. et al., 2016, *AJ*, 152, 41
- RensonP., ManfroidJ., 2009, *A&A*, 498, 961
- RiceJ. B., WehlauW. H., HolmgrenD. E., 1997, *A&A*, 326, 988
- RosénL., KochukhovO., AlecianE., NeinerC., MorinJ., WadeG. A., BinaM-IcS Collaboration, 2018, *A&A*, 613, A60
- RufenerF., 1988, Catalogue of Stars Measured in the Geneva Observatory Photometric System: 4. Geneva Observatory, Sauverny, Switzerland
- RufenerF., NicoletB., 1988, *A&A*, 206, 357
- RusomarovN., KochukhovO., RyabchikovaT., IlyinI., 2016, *A&A*, 588, A138
- RusomarovN., KochukhovO., LundinA., 2018, *A&A*, 609, A88
- RyabchikovaT. A., 2005, *Astron. Lett.*, 31, 388
- RyabchikovaT., LeoneF., KochukhovO., 2005, *A&A*, 438, 973
- RyabchikovaT., PiskunovN., KuruczR. L., StempelsH. C., HeiterU., PakhomovY., BarklemP. S., 2015, *Phys. Scr.*, 90, 054005
- SchneiderF. R. N., OhlmannS. T., PodsiadlowskiP., RöpkF. K., BalbusS. A., PakmorR., SpringelV., 2019, *Nature*, 574, 211
- ShulyakD., TsymbalV., RyabchikovaT., StützC., WeissW. W., 2004, *A&A*, 428, 993
- SikoraJ., WadeG. A., PowerJ., NeinerC., 2019a, *MNRAS*, 483, 2300
- SikoraJ., WadeG. A., PowerJ., NeinerC., 2019b, *MNRAS*, 483, 3127
- SilvesterJ., KochukhovO., WadeG. A., 2014a, *MNRAS*, 440, 182
- SilvesterJ., KochukhovO., WadeG. A., 2014b, *MNRAS*, 444, 1442
- SilvesterJ., KochukhovO., WadeG. A., 2015, *MNRAS*, 453, 2163
- SitnovaT. M., MashonkinaL. I., RyabchikovaT. A., 2013, *Astron. Lett.*, 39, 126
- SpruitH. C., 1999, *A&A*, 349, 189
- TakedaY., HondaS., 2016, *PASJ*, 68, 32
- TakedaY., SadakaneK., 1997, *PASJ*, 49, 367
- TayalS. S., ZatsarinnyO., 2016, *Phys. Rev. A*, 94, 042707
- ThompsonG. I., NandyK., JamarC., MonfilsA., HouziauxL., CarnochanD. J., WilsonR., 1978, Catalogue of Stellar Ultraviolet Fluxes: A Compilation of Absolute Stellar Fluxes Measured by the Sky Survey Telescope (S2/68) Aboard the ESRO Satellite TD-1. Science Research Council, Royal Observatory Edinburgh, Edinburgh
- TodyD., 1986, in CrawfordD. L., ed., Proc. SPIE Conf. Ser. Vol. 627, Instrumentation in Astronomy VI. SPIE, Bellingham, p. 733
- van der BlikN. S., ManfroidJ., BouchetP., 1996, *A&AS*, 119, 547
- WadeG. A., DonatiJ.-F., LandstreetJ. D., ShorlinS. L. S., 2000, *MNRAS*, 313, 851
- ZatsarinnyO., 2006, *Comput. Phys. Commun.*, 174, 273

APPENDIX A: OBSERVING LOGS, MAGNETIC FIELD, AND RADIAL VELOCITY MEASUREMENTS

Table A1. Journal of Narval spectropolarimetric observations of 45 Her. The columns give the UT and heliocentric Julian dates of mid-exposure for each observation, the S/N of Stokes V spectrum per 1.8 km s^{-1} velocity bin at 550 nm , and the rotational and orbital phases. This is followed by the Stokes V signature detection flag (DD = definite detection, MD = marginal detection, ND = no detection), the mean longitudinal magnetic field, and the radial velocity measured from the hydrogen line cores. The average uncertainty of the latter measurements is 0.52 km s^{-1} . The last column contains a running number to indicate consecutive observations averaged for the purpose of magnetic mapping. Zero in this column points to the three low S/N observations that were excluded from that analysis.

UT date	HJD	S/N	Rotational phase	Orbital phase	Stokes V detection	$\langle B_z \rangle$ (G)	V_r (km s^{-1})	N_{avg}
2007-03-14	2454173.713	980	0.000	0.587	DD	-79 ± 7	-10.4	1
2007-03-15	2454174.709	999	0.242	0.597	DD	22 ± 7	-11.2	2
2018-04-14	2458222.638	247	0.590	0.275	ND	-47 ± 28	-12.0	3
2018-04-15	2458223.518	433	0.804	0.284	ND	4 ± 18	-12.1	4
2018-04-19	2458227.646	434	0.807	0.326	ND	-15 ± 18	-11.6	5
2018-04-24	2458232.579	381	0.005	0.375	ND	-63 ± 20	-11.5	6
2018-04-25	2458233.571	318	0.246	0.385	DD	-20 ± 24	-12.2	7
2018-04-27	2458235.540	382	0.724	0.405	DD	22 ± 20	-10.4	8
2018-05-06	2458244.537	539	0.910	0.495	ND	-58 ± 14	-10.7	9
2018-05-09	2458247.513	477	0.633	0.525	DD	61 ± 15	-9.4	10
2018-05-11	2458249.539	490	0.125	0.546	ND	-22 ± 15	-10.5	11
2018-05-17	2458255.569	659	0.590	0.606	DD	54 ± 10	-9.4	12
2018-05-18	2458256.538	552	0.825	0.616	DD	-56 ± 13	-10.1	13
2018-05-18	2458257.487	655	0.056	0.626	DD	-50 ± 11	-10.3	14
2018-05-23	2458261.502	551	0.031	0.666	ND	-62 ± 13	-10.4	15
2018-06-15	2458284.530	615	0.625	0.897	DD	43 ± 11	-13.7	16
2018-06-15	2458284.548	579	0.630	0.897	DD	33 ± 12	-13.8	16
2018-06-16	2458285.570	654	0.878	0.908	MD	-45 ± 11	-15.3	17
2018-06-16	2458285.587	582	0.882	0.908	ND	-25 ± 13	-15.5	17
2018-06-17	2458286.556	523	0.117	0.918	DD	-61 ± 14	-16.0	18
2018-06-17	2458286.573	529	0.122	0.918	DD	-51 ± 14	-16.0	18
2018-06-18	2458287.522	597	0.352	0.927	DD	55 ± 11	-16.8	19
2018-06-18	2458287.540	601	0.356	0.928	DD	64 ± 11	-16.9	19
2018-06-18	2458288.475	353	0.583	0.937	DD	13 ± 20	-15.9	20
2018-06-18	2458288.493	365	0.588	0.937	ND	71 ± 19	-15.9	20
2018-06-19	2458289.413	578	0.811	0.946	DD	-36 ± 13	-17.2	21
2018-06-19	2458289.431	563	0.816	0.947	DD	-49 ± 14	-17.1	21
2018-06-24	2458293.511	541	0.807	0.988	DD	-22 ± 14	-18.9	22
2018-06-24	2458293.528	552	0.811	0.988	DD	-15 ± 14	-19.2	22
2018-06-24	2458294.490	517	0.045	0.997	MD	-72 ± 14	-19.6	23
2018-06-25	2458294.507	587	0.049	0.998	MD	-52 ± 12	-19.5	23
2018-06-25	2458295.461	586	0.281	0.007	DD	37 ± 12	-20.7	24
2018-06-25	2458295.480	488	0.285	0.007	DD	52 ± 15	-20.6	24
2018-06-30	2458300.459	72	0.495	0.057	ND	156 ± 95	-18.9	0
2018-06-30	2458300.477	277	0.499	0.058	ND	24 ± 24	-18.8	25
2018-07-09	2458309.465	647	0.683	0.148	DD	38 ± 11	-15.0	26
2018-07-17	2458317.487	641	0.631	0.228	DD	39 ± 11	-12.6	27
2018-07-22	2458322.489	709	0.846	0.279	MD	-34 ± 11	-12.5	28
2018-07-23	2458323.392	372	0.066	0.288	ND	-60 ± 20	-12.4	29
2018-07-23	2458323.489	503	0.089	0.289	DD	-88 ± 15	-12.4	29
2018-07-25	2458325.446	394	0.565	0.308	DD	103 ± 17	-11.4	30
2018-07-26	2458325.527	388	0.584	0.309	DD	38 ± 18	-11.4	30
2018-09-19	2458381.306	84	0.135	0.870	ND	-136 ± 103	-14.1	0
2018-09-23	2458385.297	473	0.104	0.910	ND	-21 ± 16	-15.5	31
2018-10-08	2458400.270	418	0.742	0.060	DD	-28 ± 18	-18.5	32
2018-10-12	2458404.263	106	0.711	0.101	ND	-5 ± 77	-16.1	0

Table A2. Journal of CES spectroscopic observations of 45 Her. The columns give the UT and mid-exposure heliocentric Julian dates for each observation, exposure times, the S/N at 550 nm, and the rotational and orbital phases.

UT date	HJD	Exposure time (s)	S/N	Rotational phase	Orbital phase
2014-05-06	2456783.569	3000	168	0.002	0.814
2014-05-07	2456785.450	4200	115	0.459	0.833
2014-05-07	2456785.500	4200	109	0.472	0.833
2014-05-11	2456789.436	4000	190	0.428	0.873
2014-05-11	2456789.486	4000	206	0.440	0.873
2014-05-12	2456789.535	4000	161	0.452	0.874
2014-05-12	2456789.580	3500	201	0.463	0.874
2014-07-10	2456849.277	4000	188	0.965	0.474
2014-07-10	2456849.325	4000	241	0.976	0.475
2014-07-10	2456849.377	4000	242	0.989	0.475
2014-07-10	2456849.425	4000	188	0.001	0.476
2014-07-10	2456849.469	3000	136	0.011	0.476
2014-07-12	2456851.287	4000	198	0.453	0.494
2014-07-12	2456851.336	4000	212	0.465	0.495
2014-07-12	2456851.384	4000	228	0.476	0.495
2014-07-12	2456851.432	4000	204	0.488	0.496
2014-07-12	2456851.480	4000	224	0.500	0.496
2014-08-09	2456879.359	4000	222	0.272	0.776
2014-08-09	2456879.408	4000	164	0.284	0.777
2014-08-10	2456880.264	4000	266	0.492	0.785
2014-08-10	2456880.314	4000	218	0.504	0.786
2014-08-10	2456880.362	4000	207	0.516	0.786
2014-08-10	2456880.410	3800	165	0.528	0.787
2014-08-11	2456881.261	4000	213	0.734	0.796
2014-08-11	2456881.309	4000	196	0.746	0.796
2014-08-11	2456881.358	4000	202	0.758	0.796
2014-08-11	2456881.405	3600	129	0.769	0.797
2015-04-28	2457140.542	1200	134	0.721	0.401
2015-04-28	2457140.558	1200	130	0.724	0.401
2015-04-28	2457140.582	1200	134	0.730	0.401
2015-04-28	2457140.597	1200	110	0.734	0.402
2015-04-28	2457140.613	1200	116	0.738	0.402
2015-06-11	2457185.443	1800	113	0.628	0.852
2015-06-11	2457185.478	1800	132	0.637	0.853
2015-06-11	2457185.501	1800	143	0.642	0.853
2015-06-12	2457185.523	1800	141	0.648	0.853
2015-06-12	2457185.546	1800	122	0.653	0.853
2015-08-01	2457236.310	2700	108	0.985	0.363
2015-08-01	2457236.381	2700	101	0.002	0.364
2015-08-01	2457236.416	2700	82	0.011	0.365
2015-08-02	2457237.368	2700	120	0.242	0.374
2015-08-02	2457237.401	2700	107	0.250	0.374
2018-05-01	2458240.472	1800	112	0.922	0.455
2018-05-01	2458240.496	1800	111	0.928	0.455
2018-05-02	2458240.520	1800	114	0.934	0.455
2018-05-02	2458240.543	1800	124	0.940	0.455
2018-05-02	2458240.566	1800	83	0.945	0.455
2018-05-02	2458240.590	1800	109	0.951	0.456

Table A3. Nightly mean radial velocities of 45 Her measured from the H α line in CES observations. The columns give the heliocentric Julian date, orbital phase, and radial velocity. The average uncertainty of these measurements is 0.71 km s^{-1} .

HJD	Orbital phase	V_r (km s^{-1})
2456783.570	0.814	−12.8
2456785.475	0.833	−12.8
2456789.510	0.873	−13.9
2456849.375	0.475	−9.5
2456851.360	0.495	−11.7
2456879.383	0.777	−11.4
2456880.338	0.786	−11.3
2456881.333	0.796	−11.3
2457140.578	0.401	−11.3
2457185.499	0.853	−12.6
2457236.369	0.364	−11.9
2457237.384	0.374	−12.1
2458240.525	0.455	−12.2

APPENDIX B: DI LINE PROFILE FITS

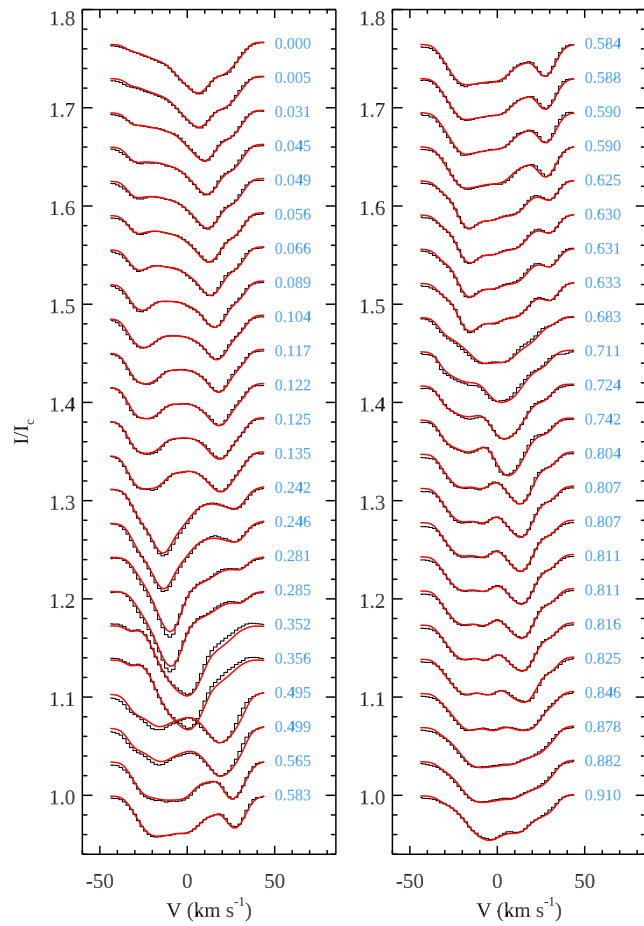


Figure B1. Comparison of the observed (black histograms) and model (red solid lines) Ti Narval Stokes I LSD profiles.

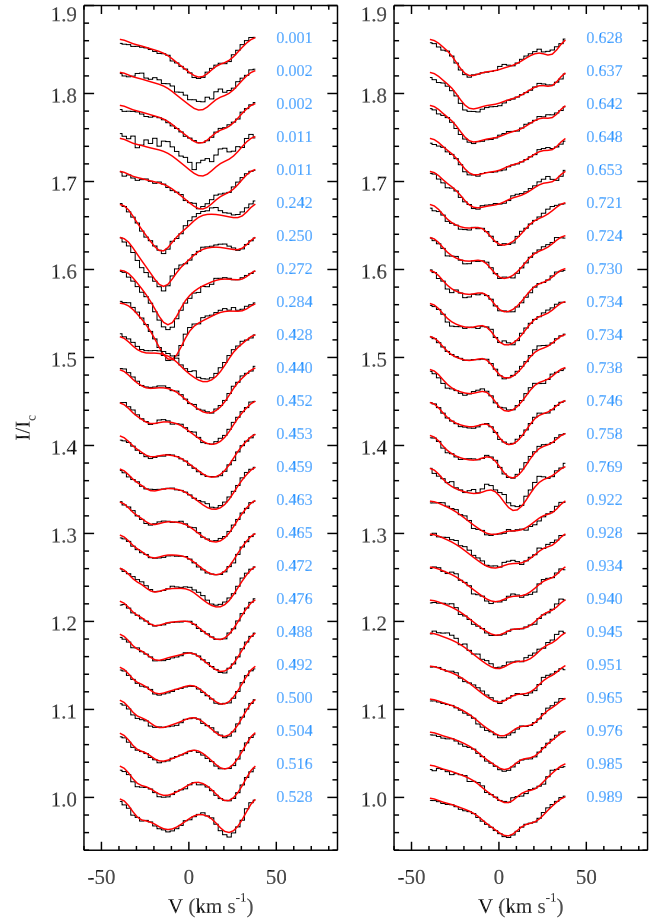


Figure B2. Same as Fig. B1 but for the Ti Stokes I LSD profiles derived from CES observations.

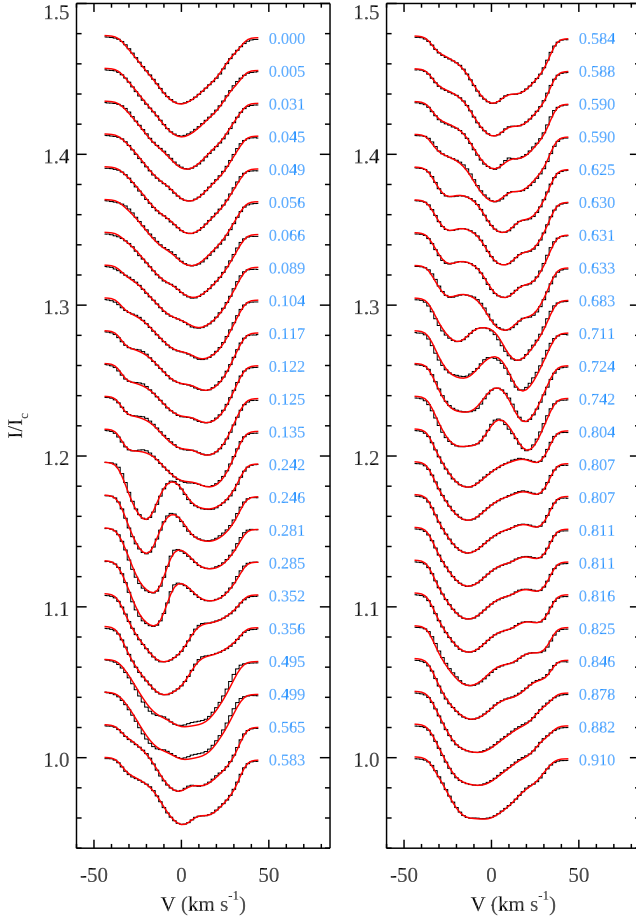


Figure B3. Same as Fig. B1, but for Cr Narval Stokes I LSD profiles.

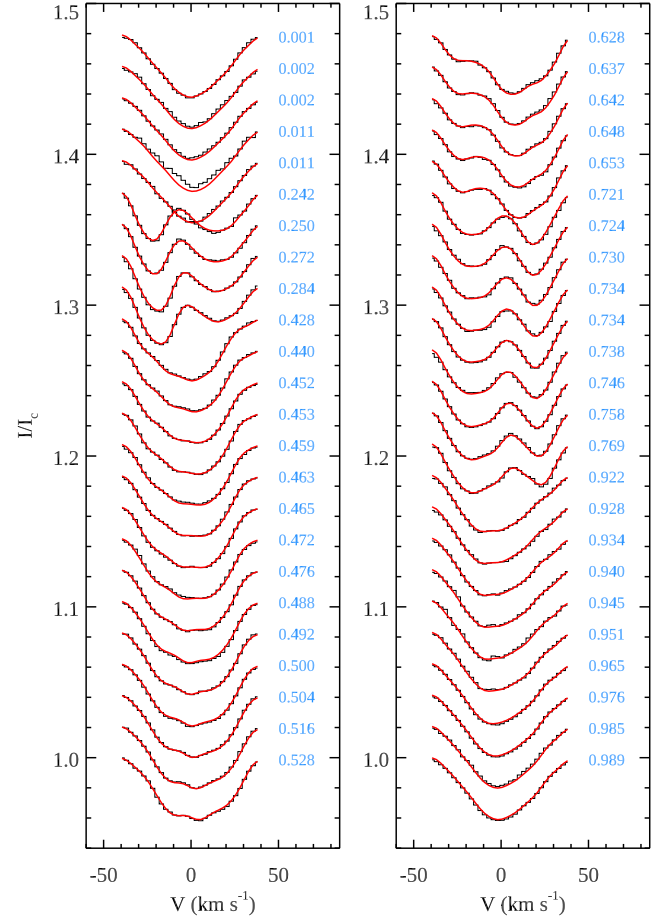


Figure B4. Same as Fig. B3 but for the Cr Stokes I LSD profiles derived from CES observations.

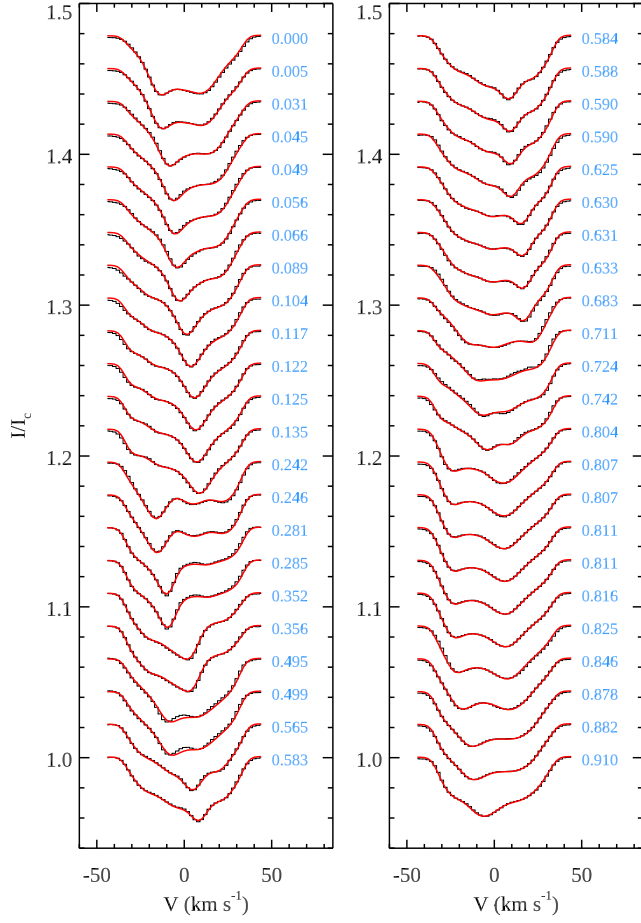


Figure B5. Same as Fig. B1, but for Fe Narval Stokes *I* LSD profiles.

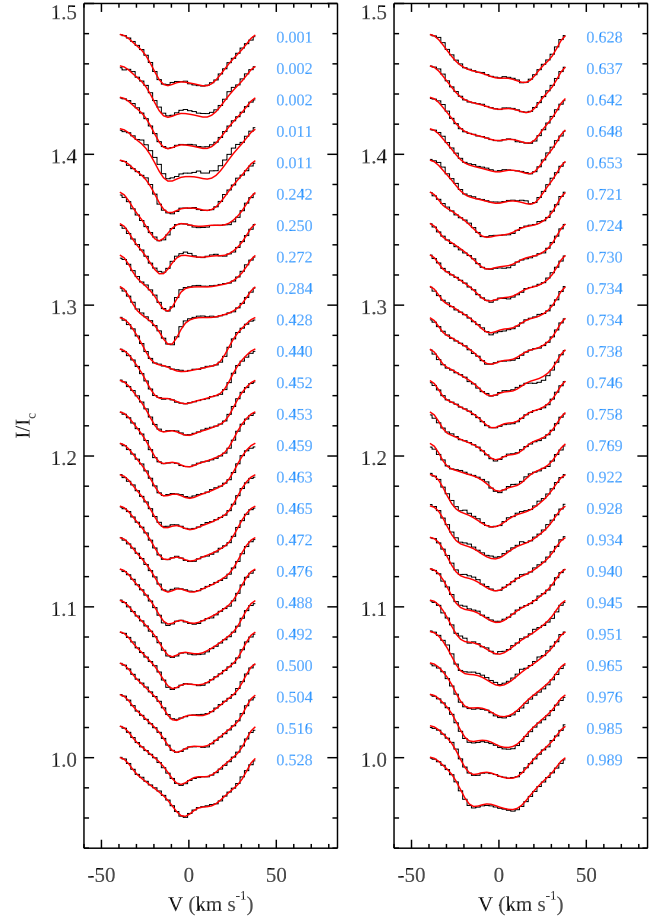


Figure B6. Same as Fig. B5 but for the Fe Stokes *I* LSD profiles derived from CES observations.

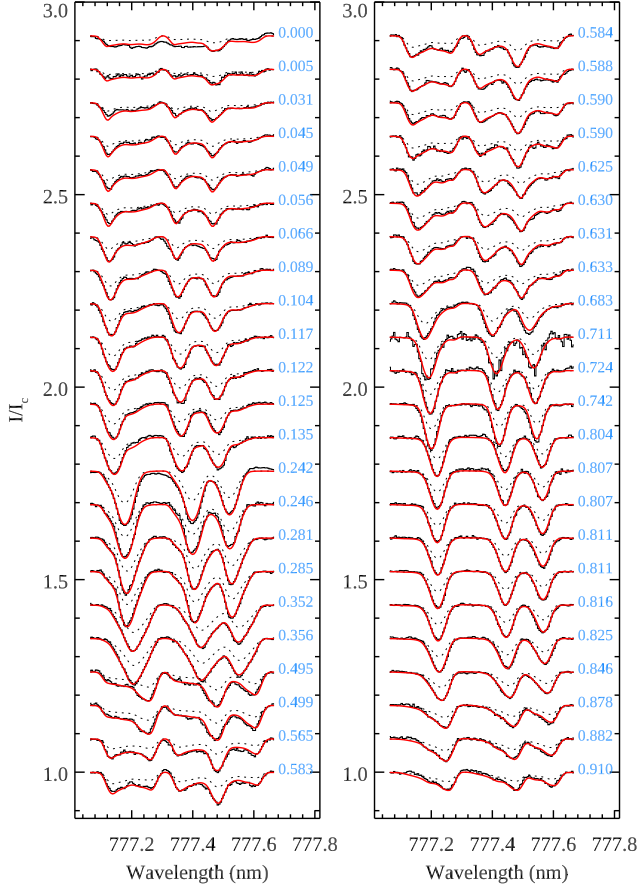


Figure B7. Same as Fig. B1, but for the O I infrared triplet in Narval spectra. Calculations accounting for NLTE effects are shown with the solid lines. The dotted lines illustrate LTE profiles for the same O surface distribution.

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